Copernicus Webinar @ Online

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Ultralight dark matter searches with laser interferometry

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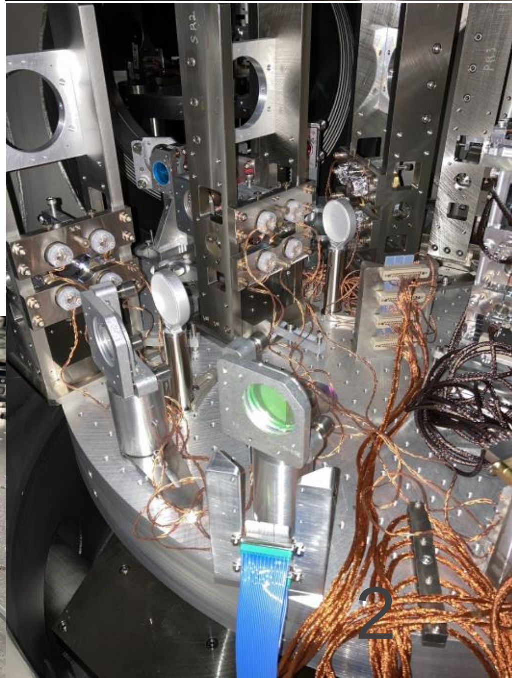
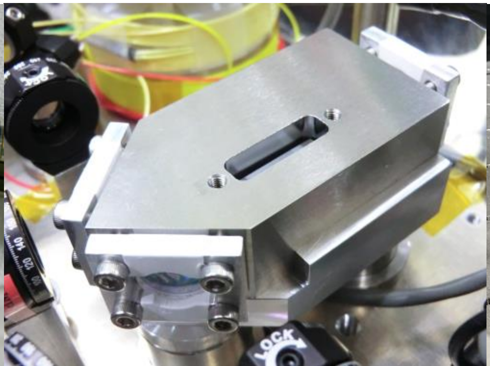
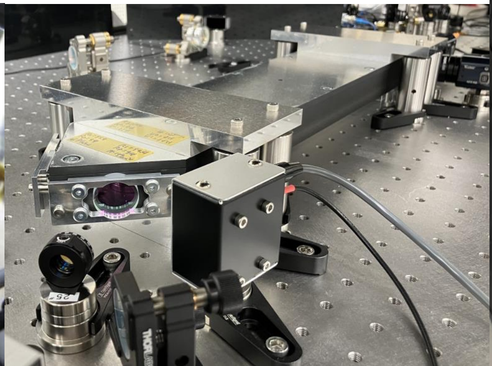
Self Introduction

- Yuta Michimura (道村唯太)
  Research Scientist at LIGO Laboratory, Caltech
- Laser interferometric gravitational wave detectors
  - Ground based: LIGO, KAGRA
  - Space based: DECIGO (SILVIA)
- Searches for new physics
  with laser interferometry
  - Lorentz invariance test
  - Optomechanics for gravity/quantum test
  - Dark matter searches etc…

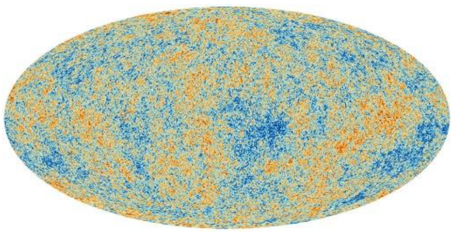






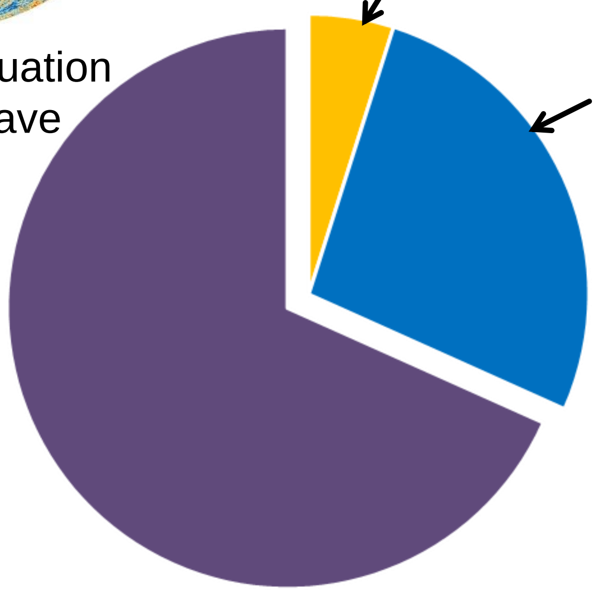
Dark Matter Mystery

- Suggested in 1930s from galaxy rotation curves
- Accounts for ~80% of all the matter of the universe
- The nature remains mystery



Temperature fluctuation of cosmic microwave background

Derive energy



Ordinary matter (~5%)
can be explained by the standard model of particle physics

Dark matter (~27%)
not included in the standard model of particle physics

All the matter

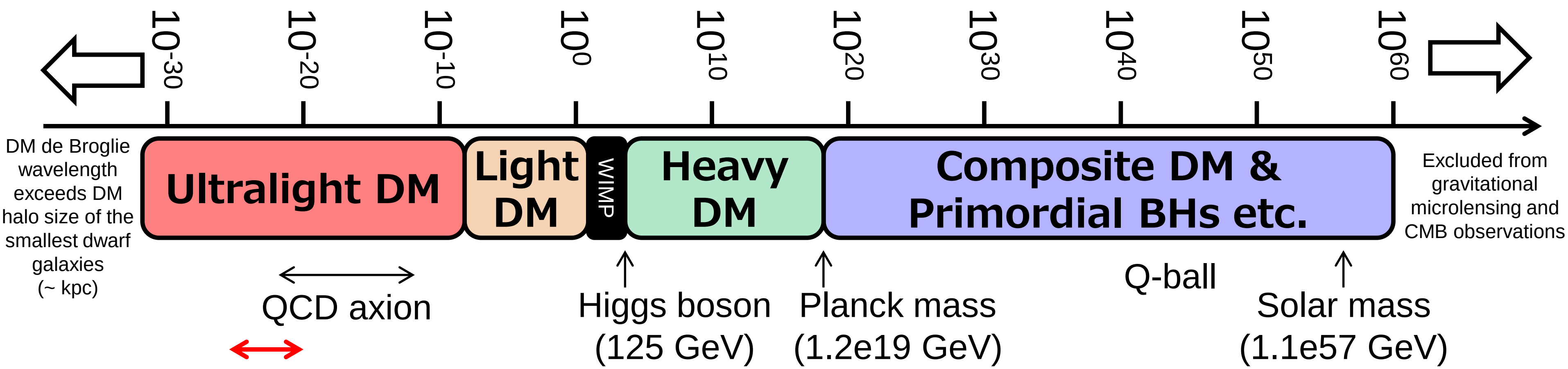
Dark energy (~68%)
drives an acceleration of the expansion of the universe

This standard model of cosmology is amazingly consistent with various observations

Dark Matter Models

- ~90 orders of magnitude
- Searches focused on WIMPs, but not detected yet
- Motivates new searches for other candidates

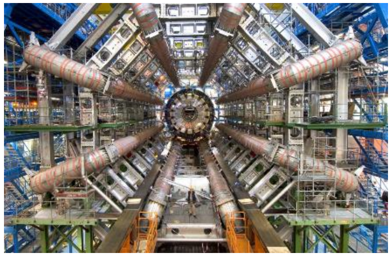
Dark Matter Mass (GeV)



2.4 Hz ~ 2.4 kHz
(1e-14 ~ 1e-11 eV)

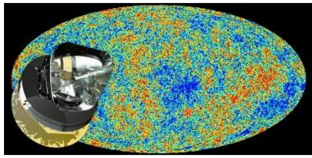
Laser Interferometry

Underground experiments (XENON1T etc.)



LHC

Subaru telescope etc.



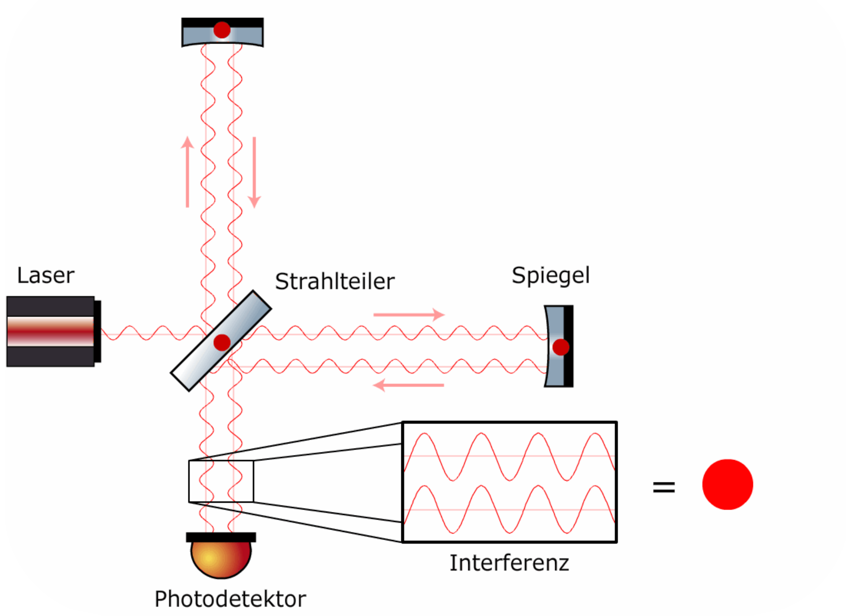
Cosmic microwave background

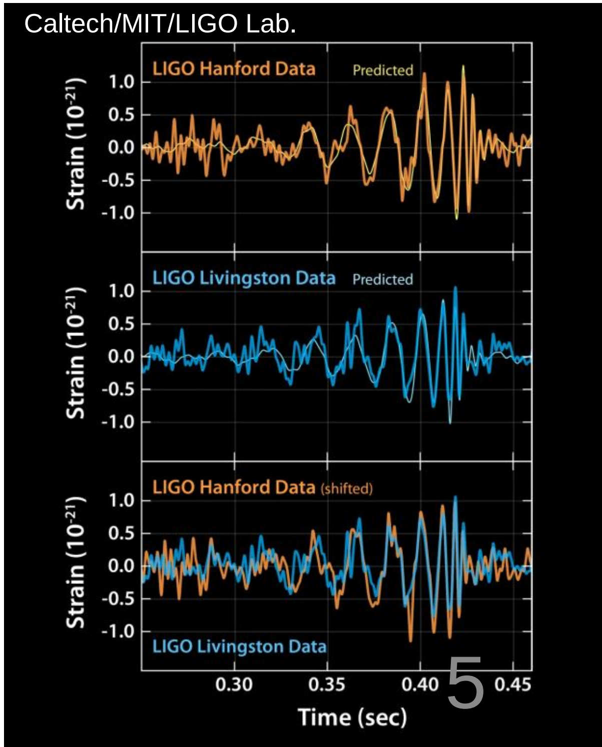
Ultralight DM with Interferometers

- Bosonic ultralight field (<~1 eV) are well-motivated from cosmology
- Behaves as classical waves

$$f = 242\ \mathrm{Hz} \left(\frac{m_{\mathrm{DM}}}{10^{-12}\ \mathrm{eV}} \right)$$

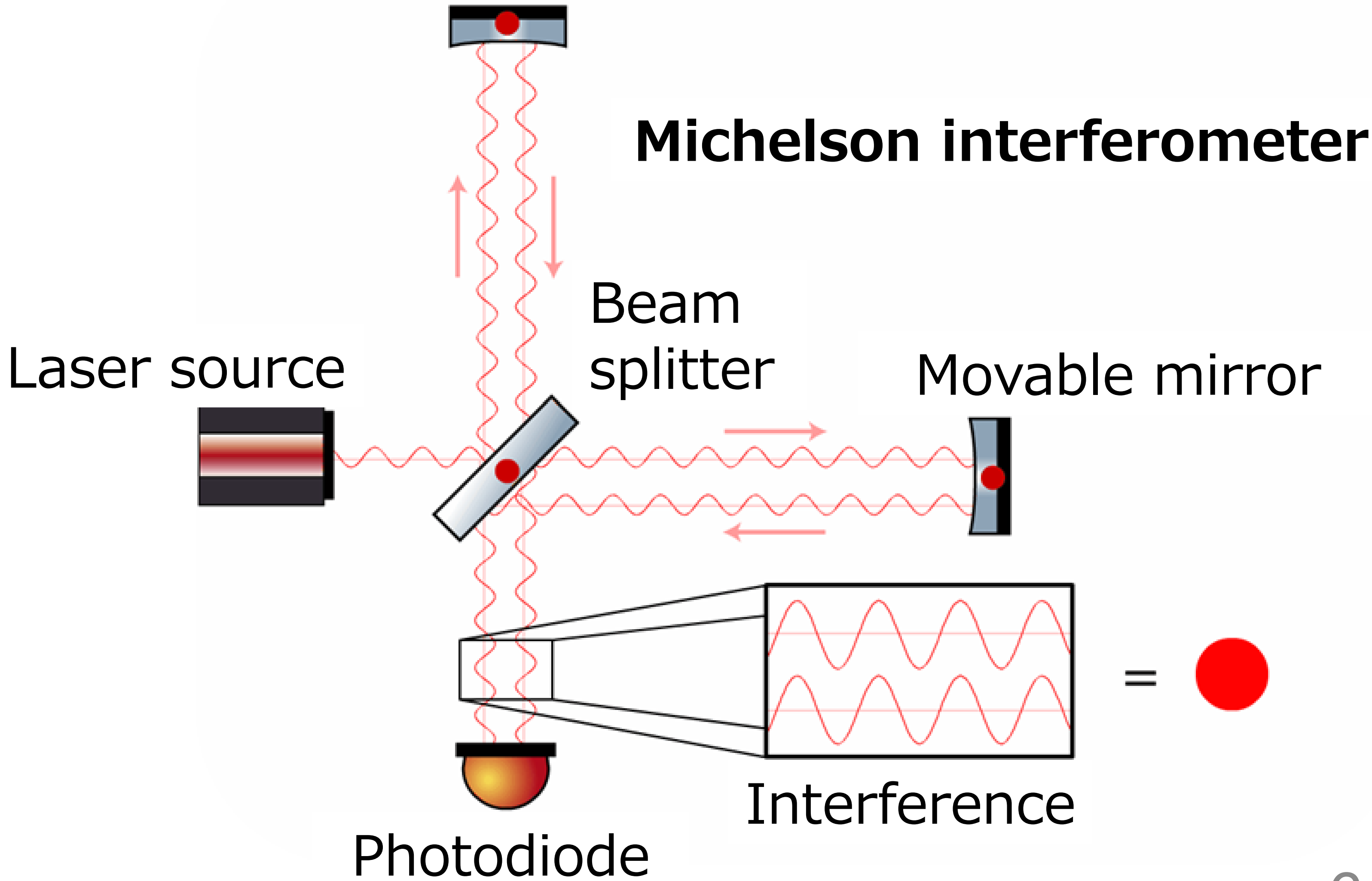
- Laser interferometers are sensitive to such oscillating changes





Laser Interferometry

- measures differential arm length change



Laser Interferometry

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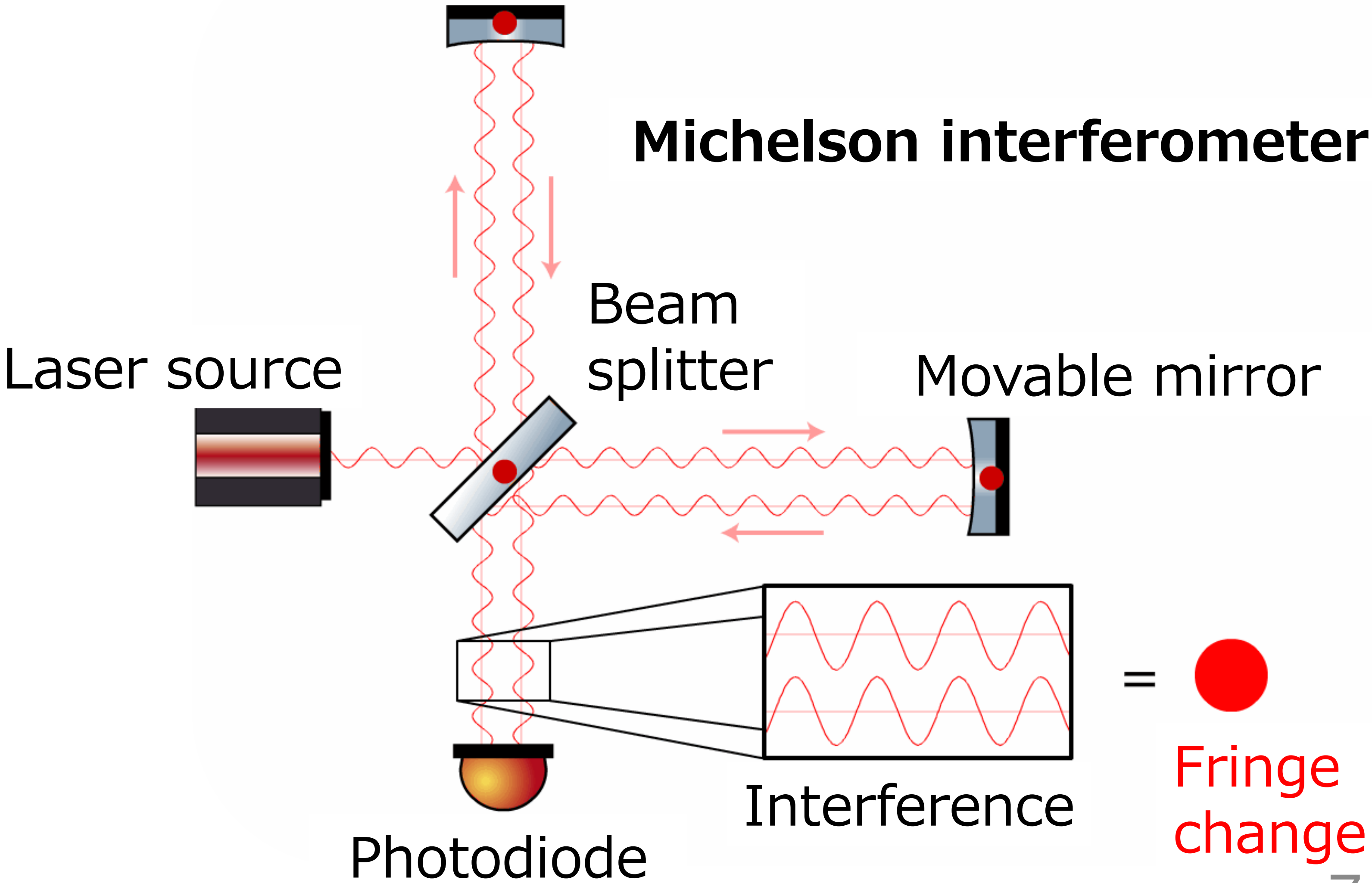

Laser Interferometry

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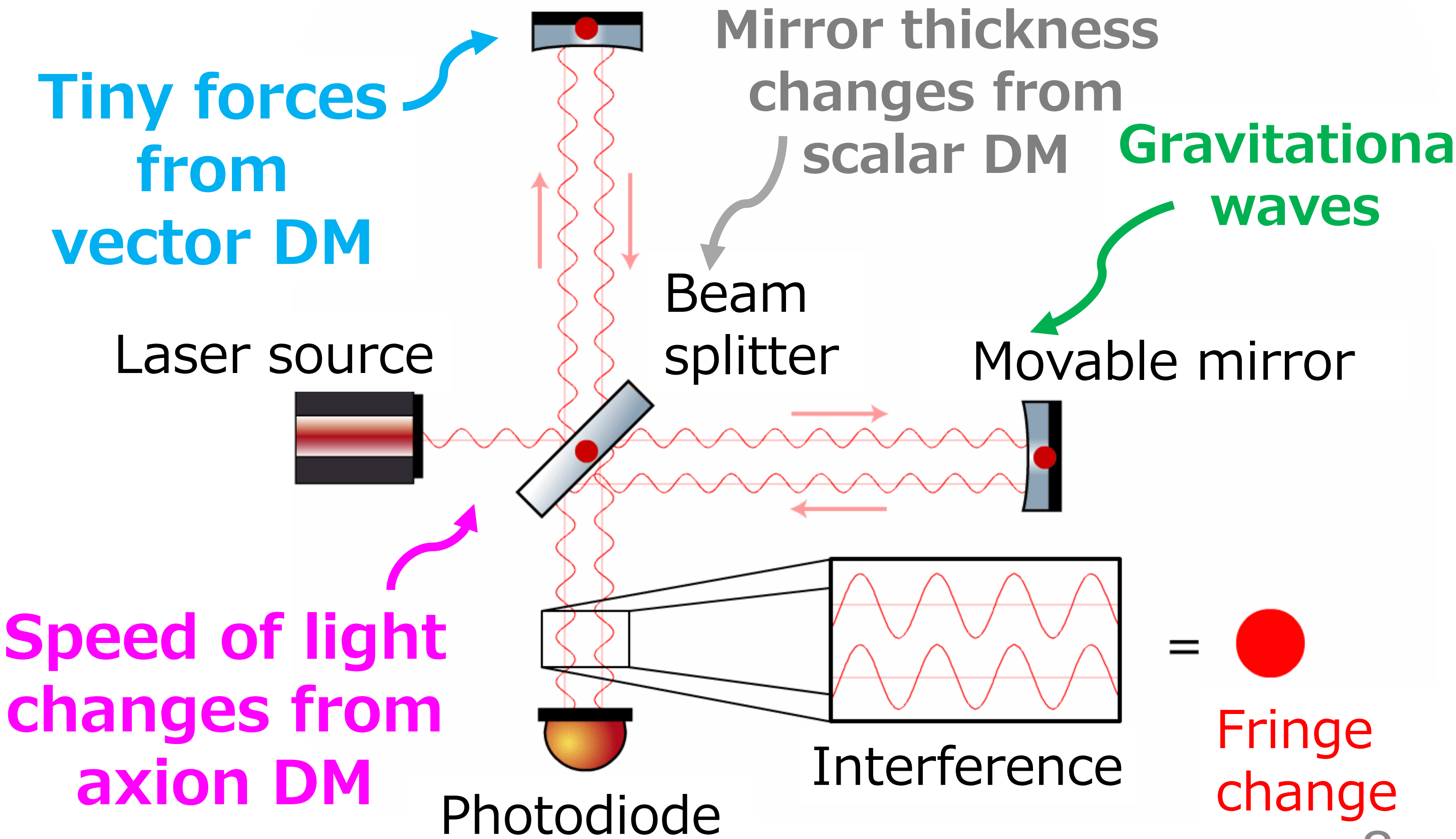


Recent Proposals and Searches

- **$U(1)_B$ or $U(1)_{B\text{-}L}$ gauge bosons(vector field)**
  - P. W. Graham+, PRD **93**, 075029 (2016)
  - A. Pierce+, PRL **121**, 061102 (2018)
  - H-K Guo+, Commun. Phys. **2**, 155 (2019) **LIGO O1 data analysis**
  - Y. Michimura, T. Fujita, S. Morisaki, H. Nakatsuka, I. Obata, PRD **102**, 102001 (2020)
  - D. Carmey+, New J. Phys. **23**, 023041 (2021)
  - J. Manley+, PRL **126**, 061301 (2021)
  - S. Morisaki, T. Fujita, Y. Michimura, H. Nakatsuka, I. Obata, PRD **103**, L051702 (2021)
  - LIGO-Virgo-KAGRA Collaboration, PRD **105**, 063030 (2022) **LIGO/Virgo O3 data analysis**
- **Scalar bosons**
  - Y. V. Stadnik & V. V. Flambaum, PRL **114**, 161301 (2015)
  - Y. V. Stadnik & V. V. Flambaum, PRA **93**, 063630 (2016)
  - A. A. Geraci+, PRL **123**, 031304 (2019)
  - H. Grote & Y. V. Stadnik, PRR **1**, 033187 (2019)
  - S. Morisaki & T. Suyama, PRD **100**, 123512 (2019)
  - C. Kennedy+, PRL **125**, 201302 (2020)
  - E. Savalle+, PRL 126, 051301 (2021)
  - S. M. Vermeulen+, Nature **600**, 424 (2021) **GEO600 data analysis**
  - K. Fukusumi, S. Morisaki, T. Suyama, arXiv:2303.13088 **LIGO/Virgo O3 data analysis**

Many recent proposals

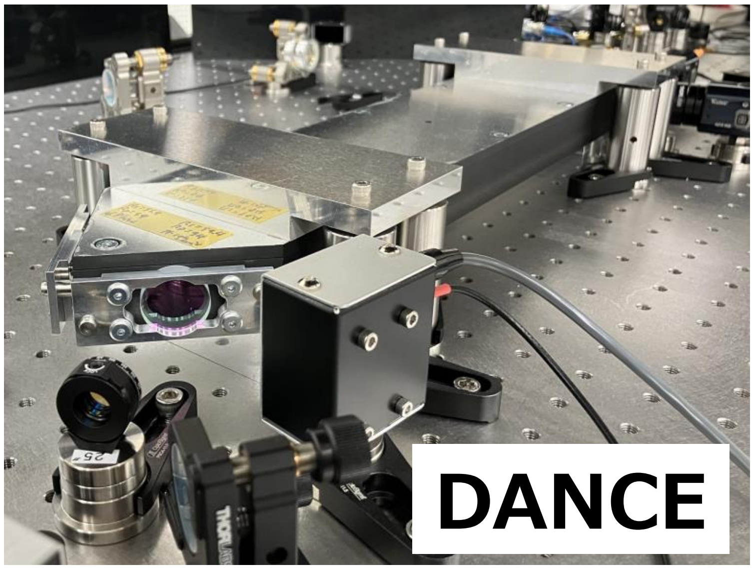
First searches with real data from GW detectors already done for gauge bosons and scalar bosons.

- **Axion & axion-like particles (ALPs)**
  - W. DeRocco & A. Hook, PRD **98**, 035021 (2018)
  - I. Obata, T. Fujita, Y. Michimura, PRL **121**, 161301 (2018)
  - H. Liu+, PRD **100**, 023548 (2019)
  - K. Nagano, T. Fujita, Y. Michimura, I. Obata, PRL **123**, 111301 (2019)
  - D. Martynov & H. Miao, PRD **101**, 095034 (2020)
  - K. Nagano, H. Nakatsuka, S. Morisaki, T. Fujita, Y. Michimura, I. Obata, PRD **104**, 062008 (2021)
  - Y. Oshima+, arXiv:2303.035947 **DANCE first result**

Not exhaustive.
The ones which require magnetic fields are not listed.

Our Projects

- Use both table-top optical cavities and large-scale laser interferometric gravitational wave detectors



Ultralight DM

Broad band

PRL **121**, 161301 (2018)

Narrow band

PRL **123**, 111301 (2019)

PRD **104**, 062008 (2021)

+ **Polarization measurement**

Axion like particles

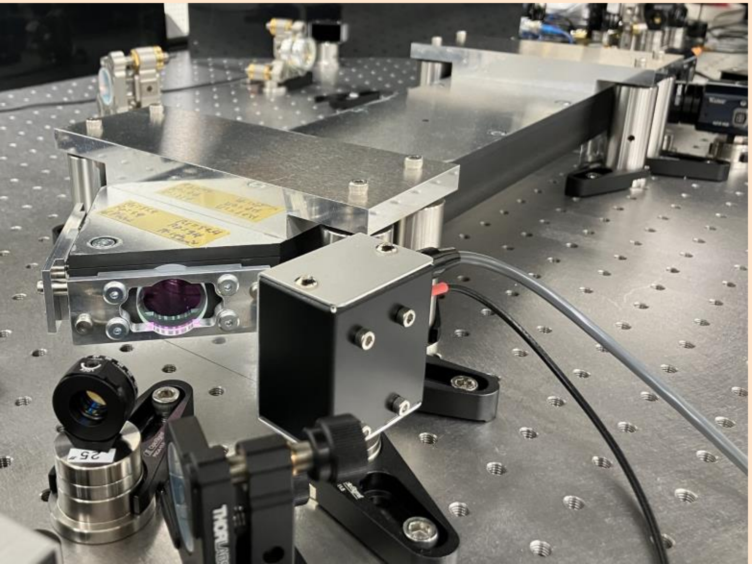
Vector bosons

PRD **102**, 102001 (2020)

PRD **103**, L051702 (2021)

Contents

- Axion dark matter search
  with table-top optical ring cavity
- Axion dark matter search
  with gravitational wave detectors
- Vector dark matter search
  with gravitational wave detectors



Contents

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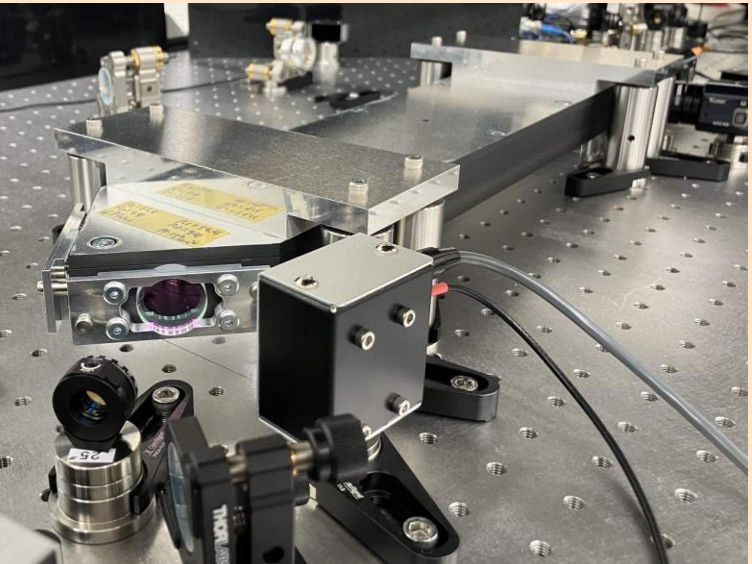
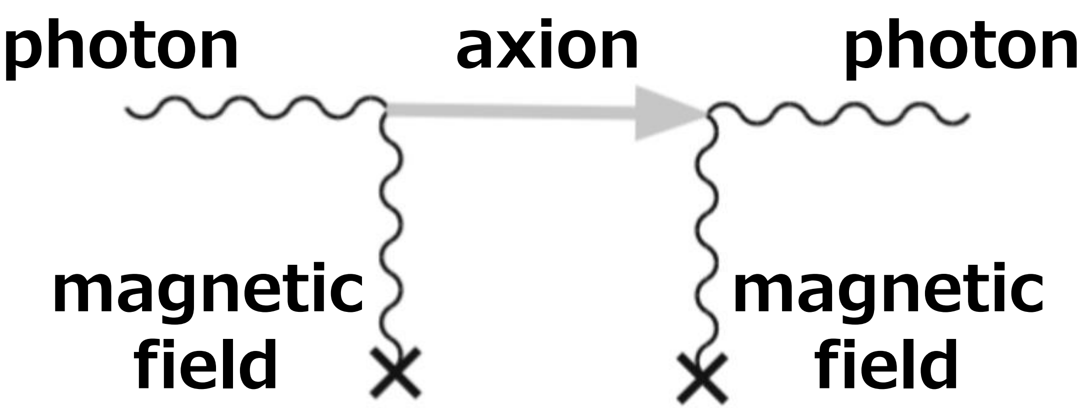
- Axion dark matter search with gravitational wave detectors
- Vector dark matter search with gravitational wave detectors



Axion and Axion-Like Particles

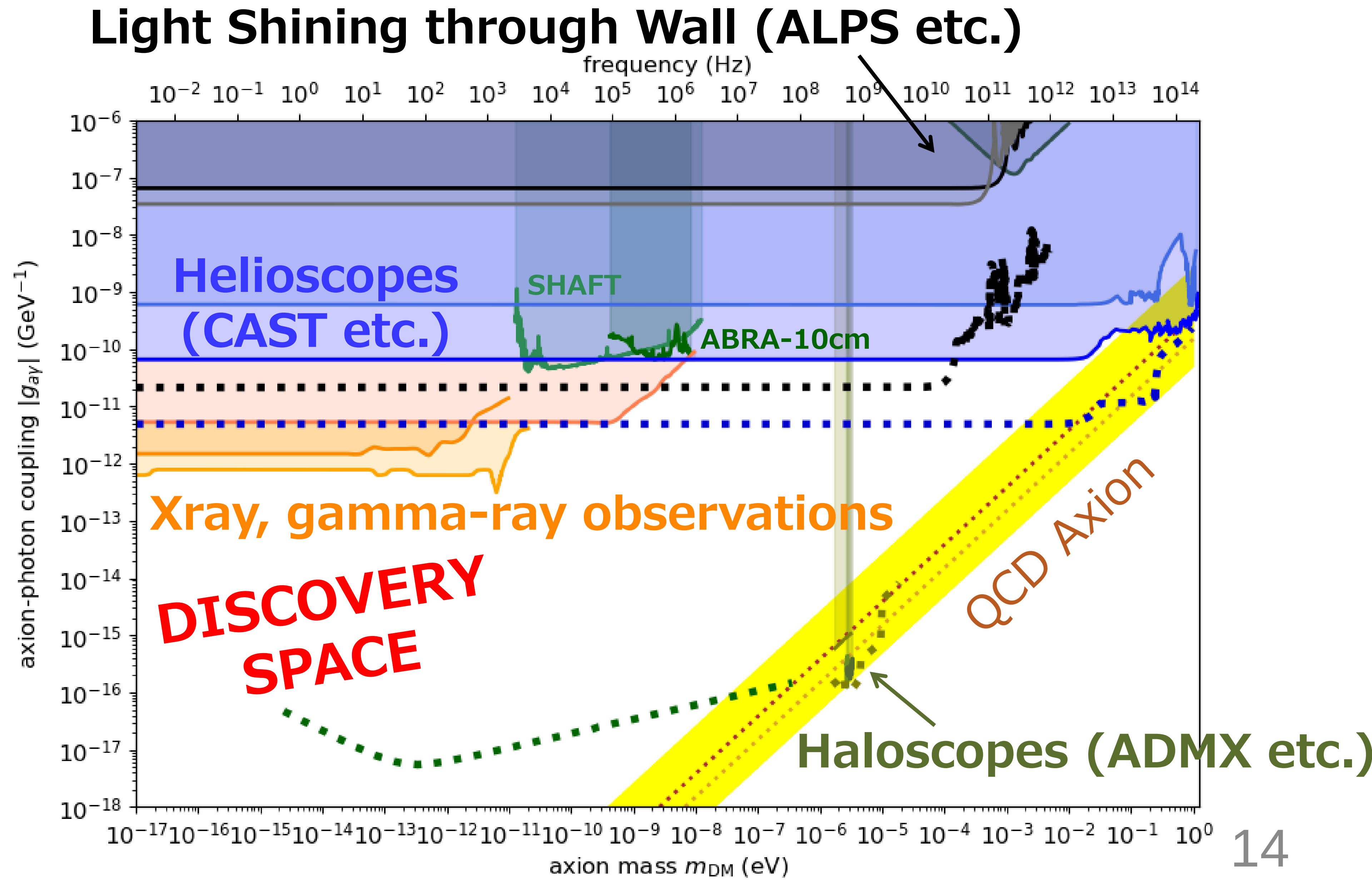
- Pseudo-scalar particle originally introduced to solve strong CP problem (QCD axion)
- Various axion-like particles (ALPs) predicted by string theory and supergravity
- Many experiments to search for ALPs through axion-photon coupling

  Especially by using magnetic fields



Primakoff effect

Previous Searches



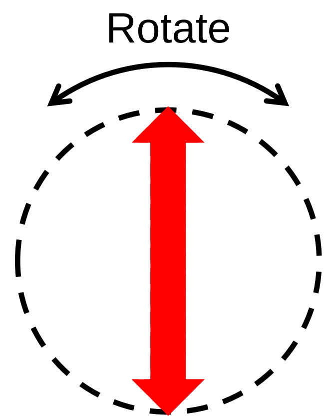
Polarization Modulation from Axions

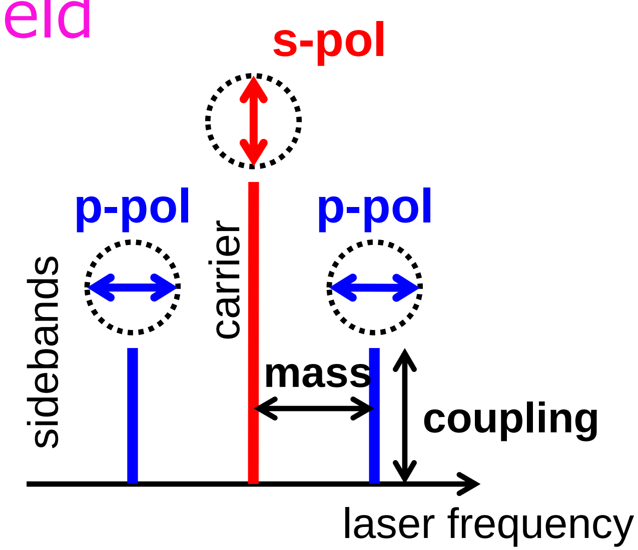
- Axion-photon coupling ($\frac{g_{a\gamma}}{4} a F_{\mu\nu} \tilde{F}^{\mu\nu}$) gives different phase velocity between left-handed and right-handed circular polarizations

$$c_{\mathrm{L/R}} = \sqrt{1 \pm \frac{g_{a\gamma} a_0 m_a}{k} \sin(m_a t + \delta_\tau)}$$

coupling constant   axion field   axion mass

- Linear polarization will be modulated
  p-pol sidebands will be generated from s-pol
- Search can be done without magnetic field





Optical Cavity to Amplify the Signal

- Polarization rotation is small for short optical path

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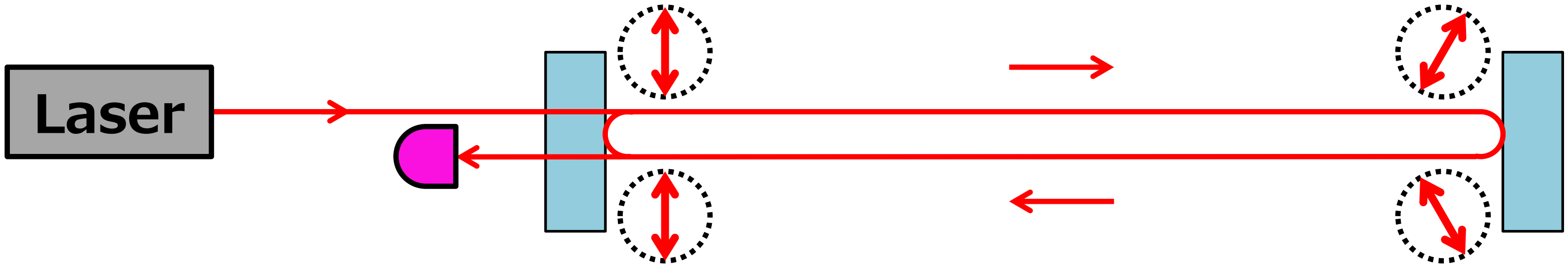
- Optical cavities can increase the optical path, but the polarization is flipped by mirror reflections



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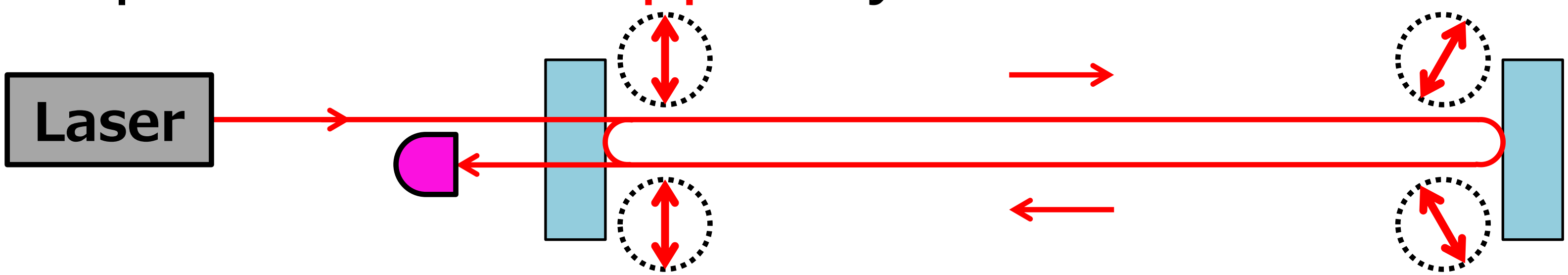
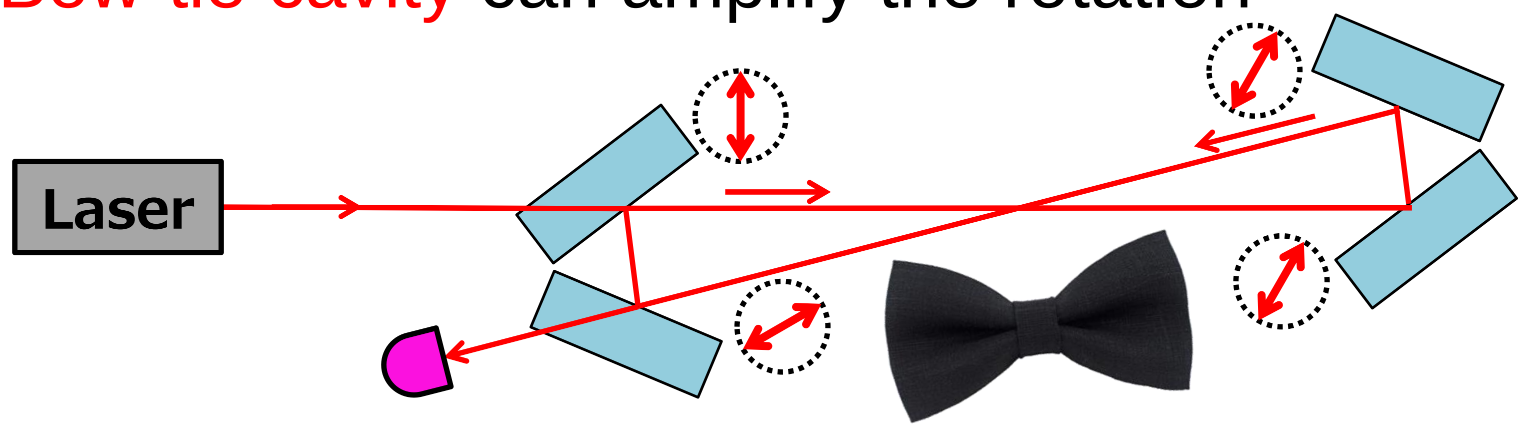


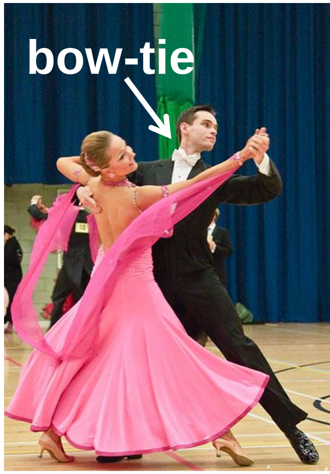
- Bow-tie cavity can amplify the rotation

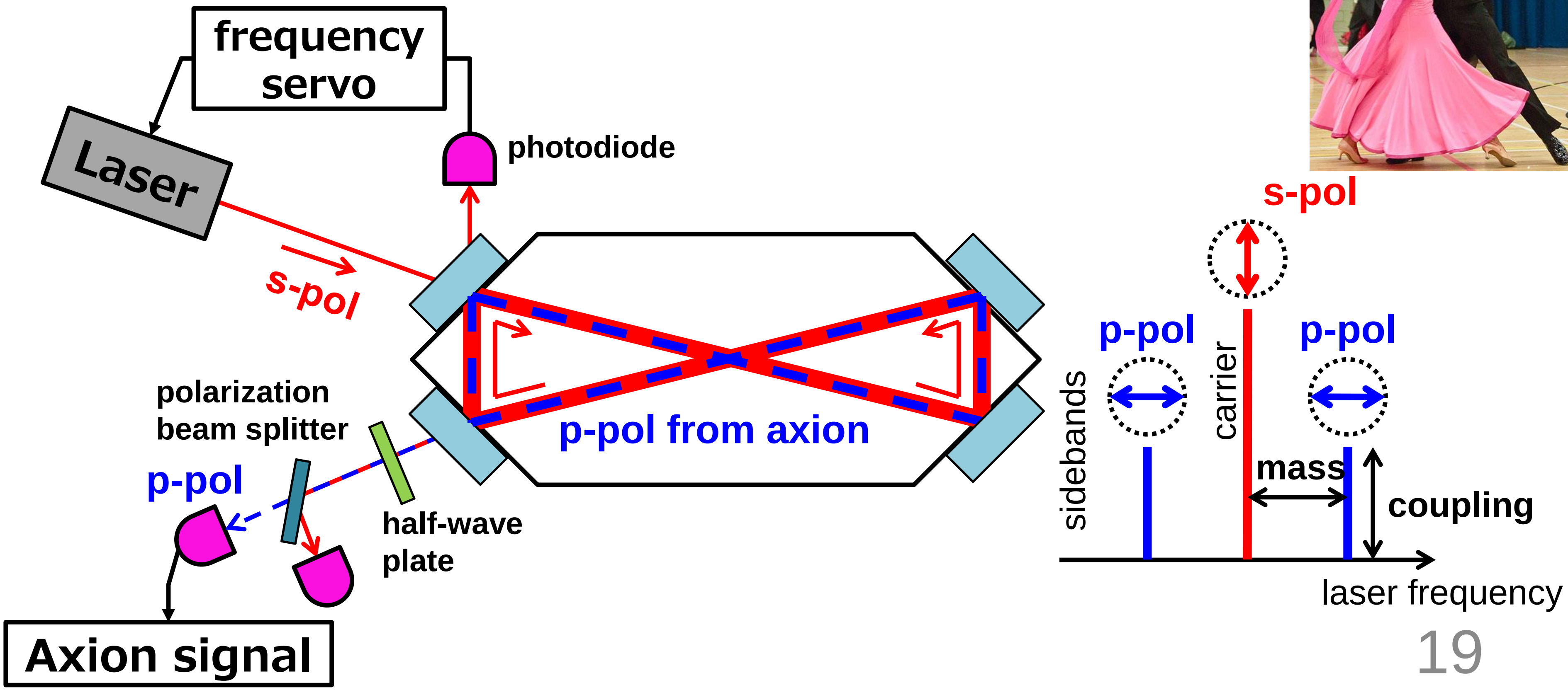


DANCE Setup

Dark matter Axion search with riNg Cavity Experiment

- Look for amount of modulated p-pol generation in each frequency

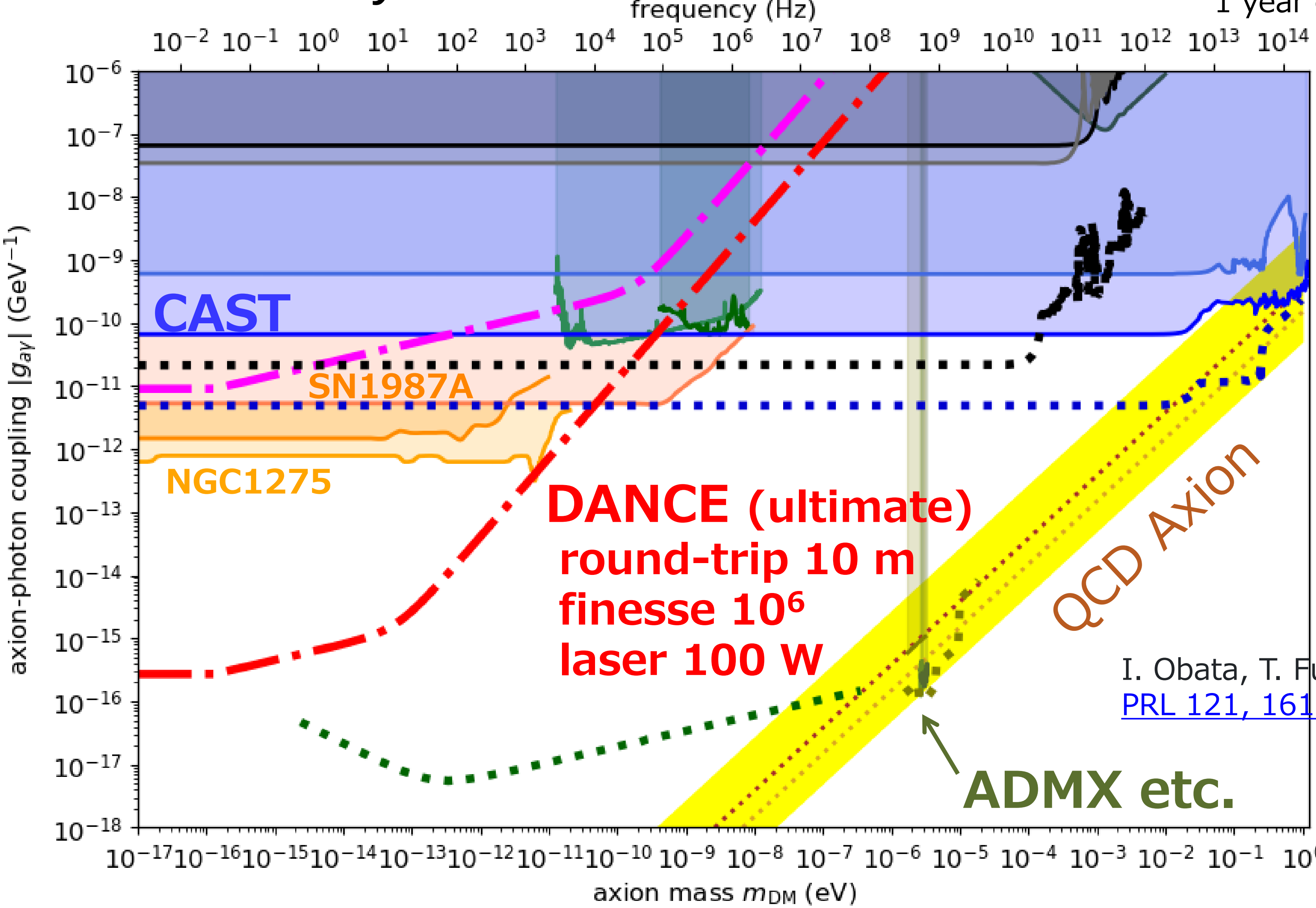




Sensitivity of DANCE

- Sensitivity better than CAST limit

* Shot noise limited
1 year observation



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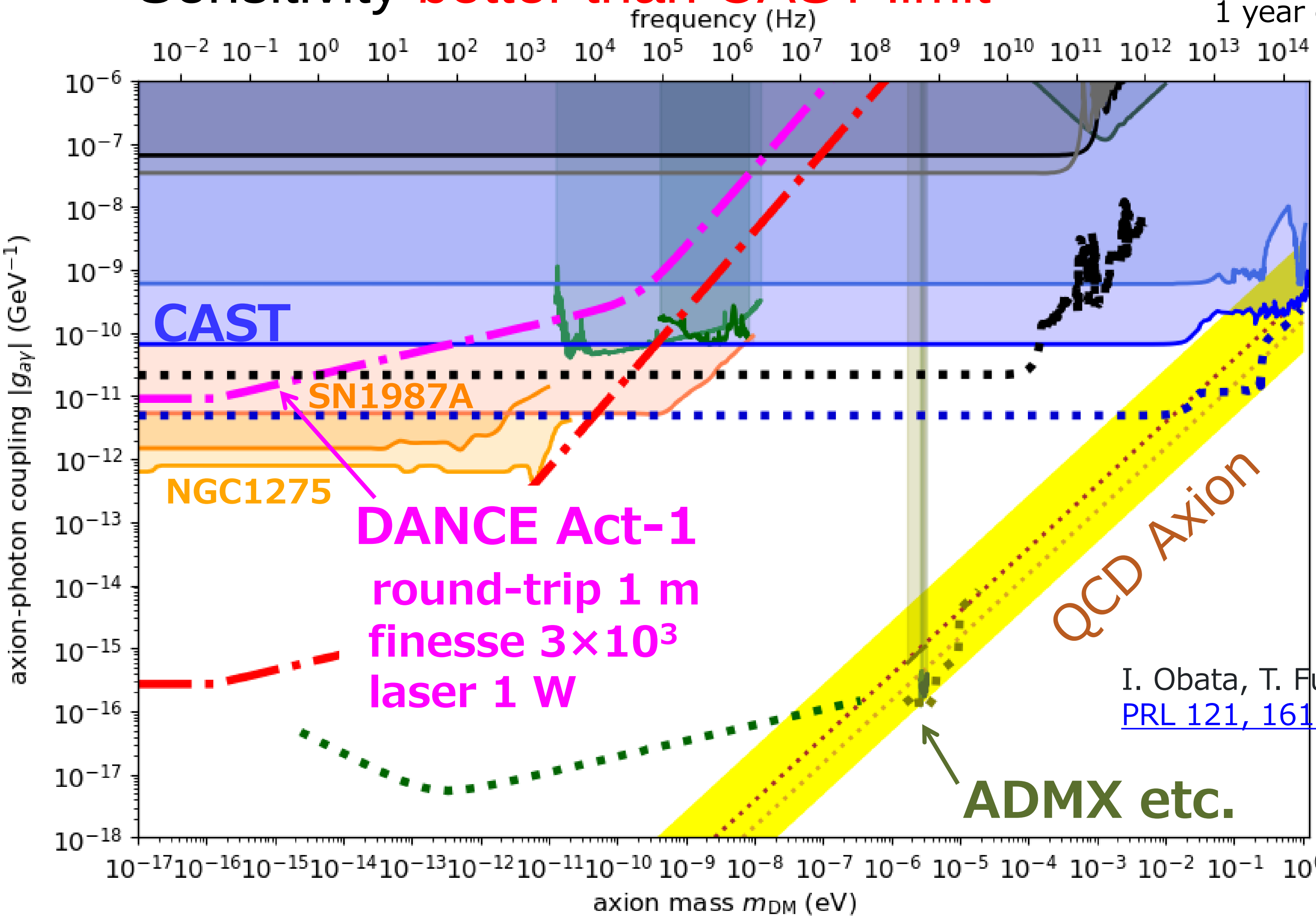
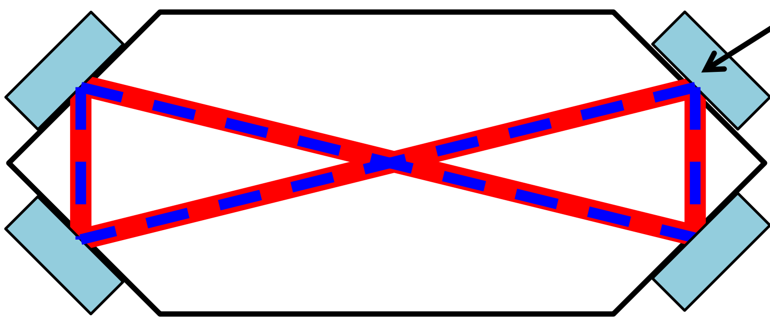


I. Obata, T. Fujita, YM,
[PRL 121, 161301 (2018)]

Status of DANCE Act-1

- Started in 2019
- After reassembly of the optics by several times and installation of digital servo system for long runs, first 12-day observation was achieved in May 2021
  - Issue: s-pol and p-pol do not resonate simultaneously
    Due to phase difference in mirror reflections
- Designed an auxiliary cavity, and achieved simultaneous resonance for the first time in November 2021

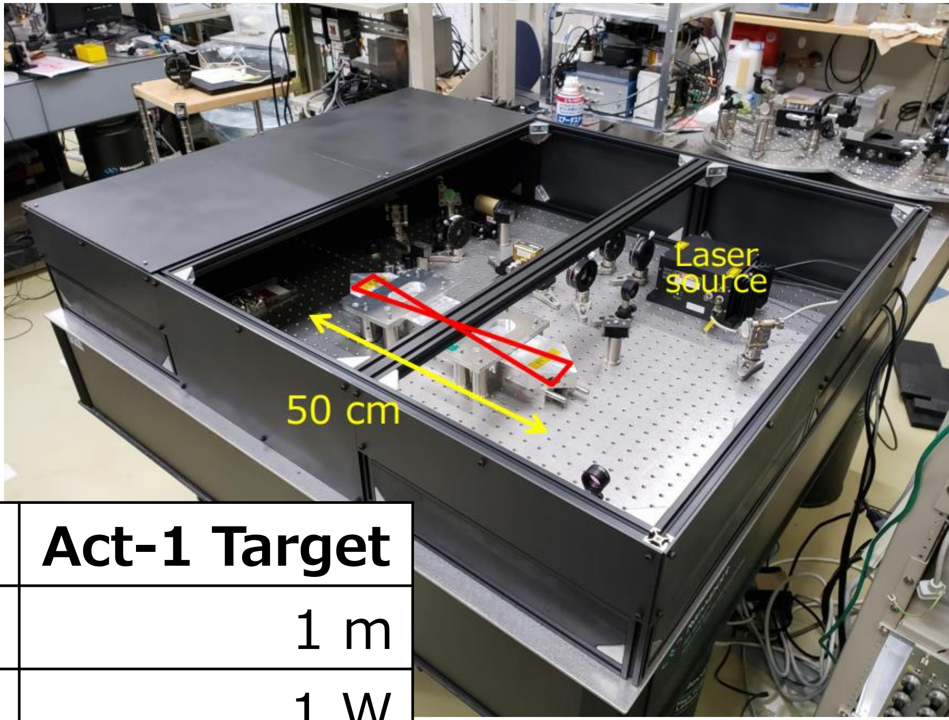


s-pol and p-pol obtain different phase on mirror reflections at non-zero incident angle
→ results in resonant frequency difference

Y. Oshima+, arXiv:2105.06252
H. Fujimoto+, arXiv:2105.08347
Y. Oshima+, JPCS **2156**, 012042 (2021)
H. Fujimoto+, JPCS **2156**, 012182 (2021)

First Observing Run in May 2021

- Same scale as Act-1 target
- 12-day test run from May 8th to 30th

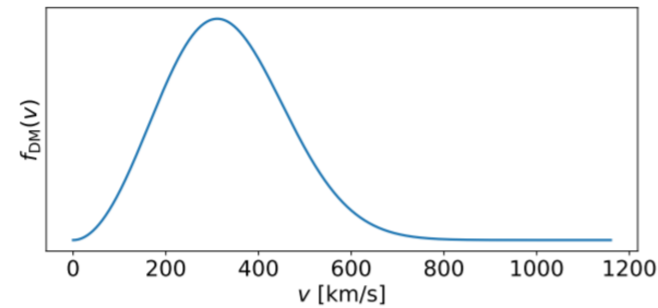


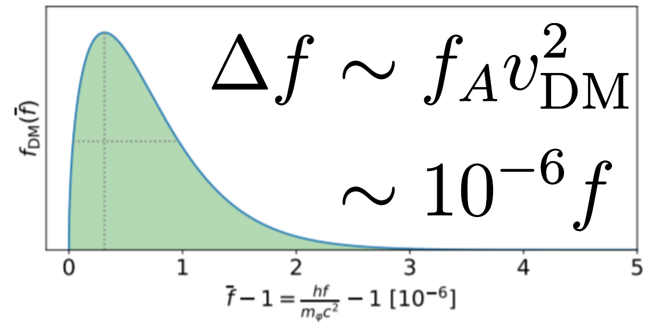
Y. Oshima+, [arXiv:2303.03594](https://arxiv.org/abs/2303.03594)

	May 2021	**Act-1 Target**
Round-trip length	1 m	1 m
Input power	242(12) mW (Source: 0.5 W)	1 W
Finesse (for carrier)	$2.85(5)\times10^3$ s-pol	3×10^3
Finesse (for sidebands)	195(3) p-pol	3×10^3
s/p-pol resonant freq. difference	2.52(2) MHz	0 Hz

Data Analysis Pipeline

E. Savalle+,
PRL 126, 051301 (2021)





- Nearly monochromatic signal

$$\omega_i = m_a \left(1 + \frac{v_i^2}{2}\right)$$

- Stack the spectra in this frequency region to calculate SNR

$$\rho = \sum \frac{4|\tilde{d}(f_k)|^2}{T_{\rm obs} S_n(f_k)}$$

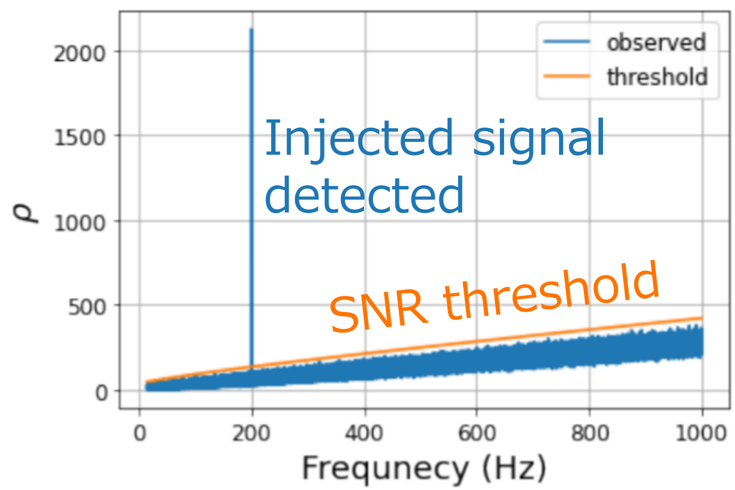
Data

PSD

$m_A \leq 2\pi f_k \leq m_A(1 + \kappa v_{\rm DM}^2)$

Obs. time

- Detection threshold determined assuming ρ follows χ^2 distribution (=assuming Gaussian noise)



- From ρ, calculate 95% upper limit on coupling constant
- Applied the pipeline to mock data for verification

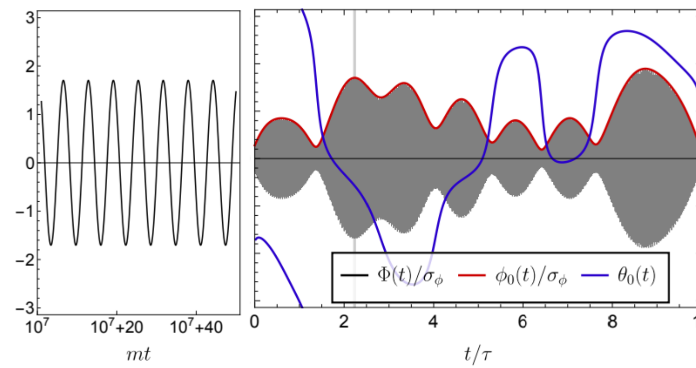
Stochastic Nature of DM Signal

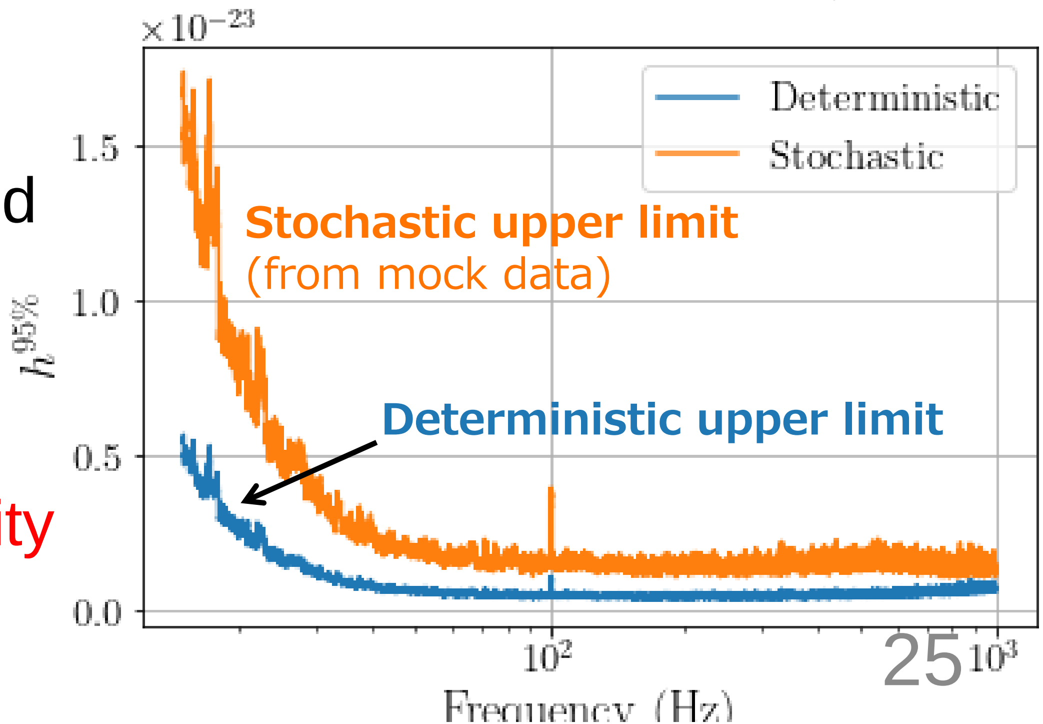
- DM signal is from superposition of many waves with various momentum, phase and polarization
- The amplitude fluctuates at the time scale of

$$\tau = 2\pi/(m_a v_{\rm DM}^2)$$

- At low frequencies, DM signal could be too small by chance and elude detection
- Method to calculate upper limit taking into account this stochasticity developed

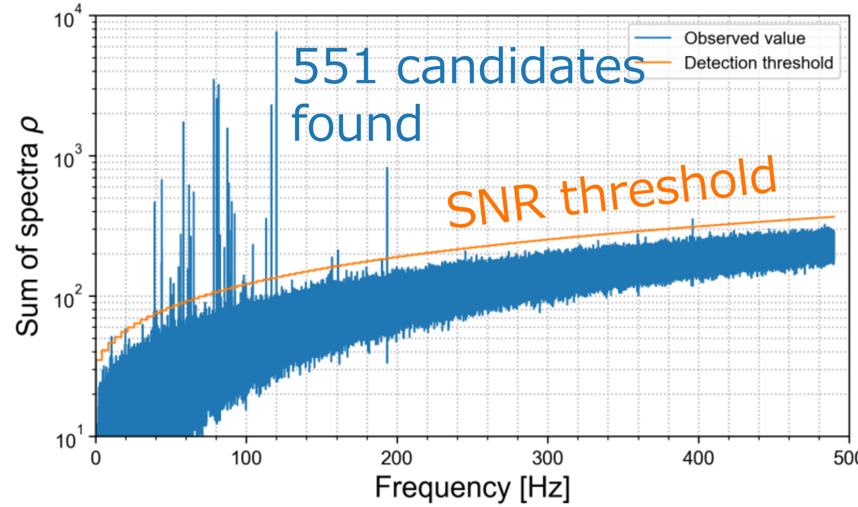
H. Nakatsuka+, arXiv:2205.02960

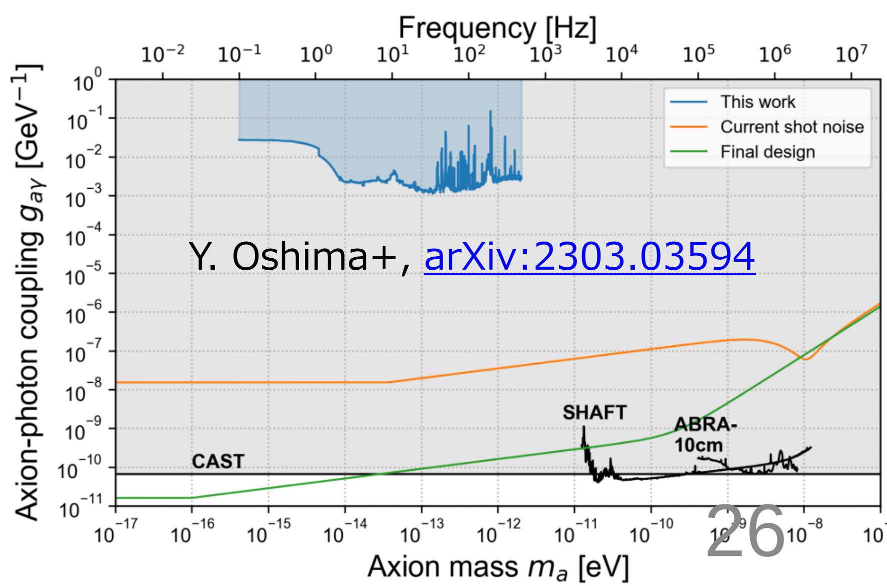




First Data Analysis Results

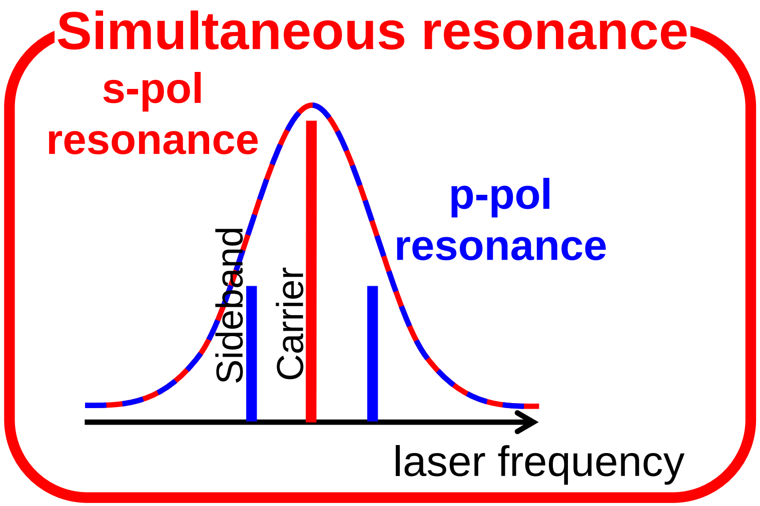
- Used 24-hour data from 12-day run
- 551 candidates found from initial analysis
- Veto analysis
  - Consistency veto
    (Frequency should be the same for different set of 24-hour data)
  - Q-factor veto
    (DM signal must have Q of 10^6)
  - Remaining 7 candidates (all multiples of ~40 Hz) are also found in laser frequency control, and thus rejected
- Placed upper limits

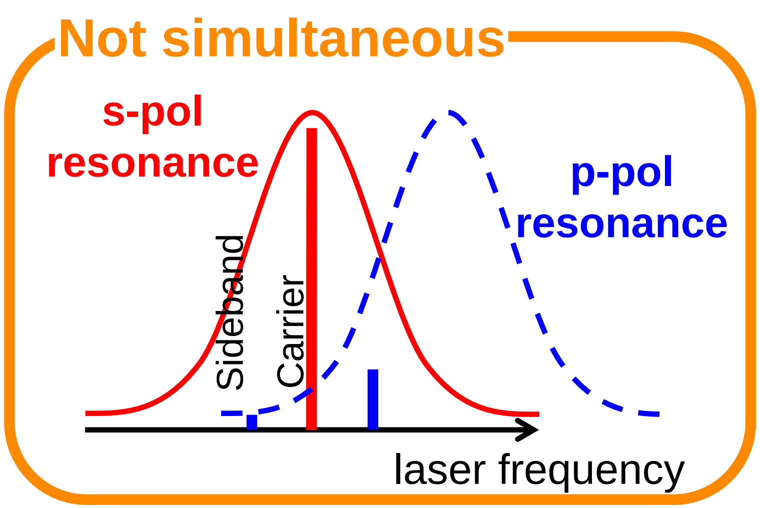


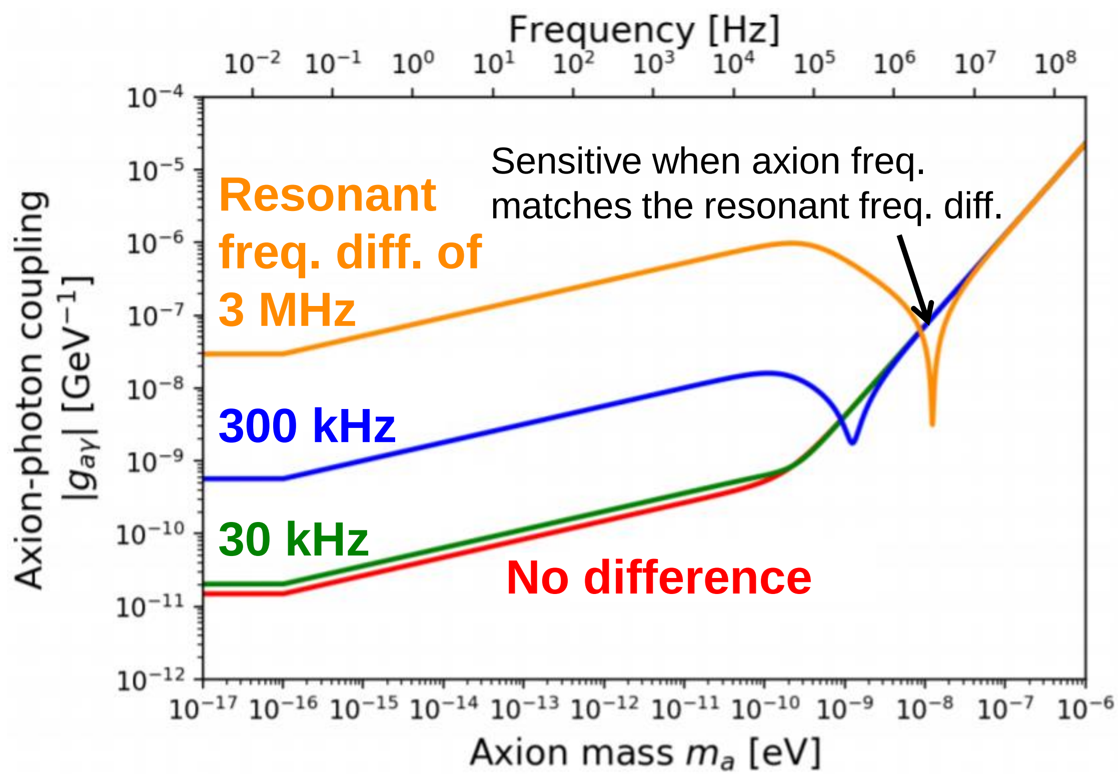


Simultaneous Resonance

- Carrier pol and sideband pol needs to be enhanced simultaneously for improving the sensitivity



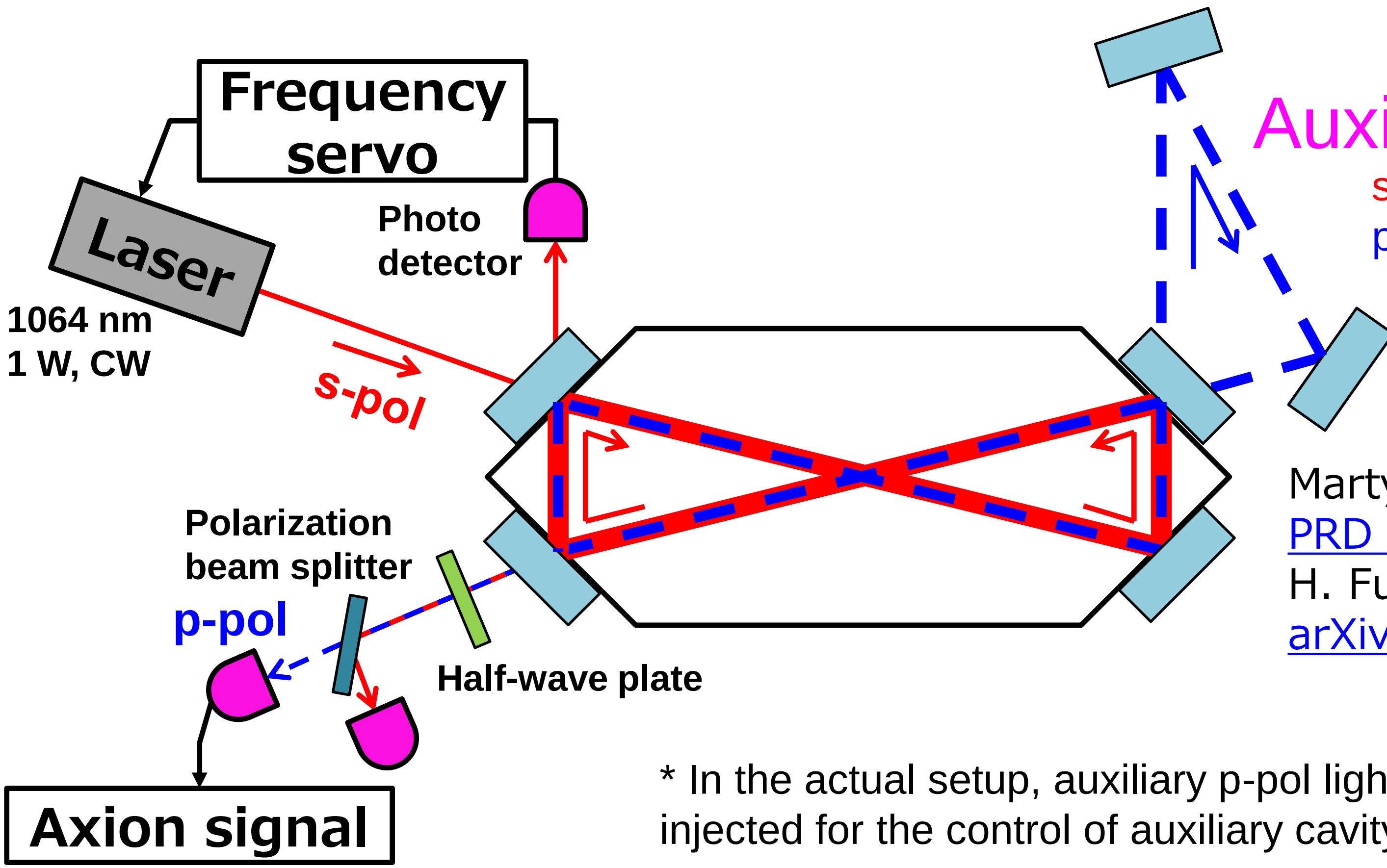




Plot by Y. Oshima & H. Fujimoto

Auxiliary Cavity as Solution

- Make resonant condition for auxiliary cavity different between s/p-pol to make reflected phase different
- This compensates phase difference in the main cavity



Martynov & Miao
PRD 101, 095034 (2020)
H. Fujimoto+,
arXiv:2110.12023

* In the actual setup, auxiliary p-pol light is also injected for the control of auxiliary cavity length

Updated Setup

- New lab prepared
- New 2W laser source obtained
  (previously, 0.5W laser source)
- Installed an auxiliary cavity

Auxiliary cavity

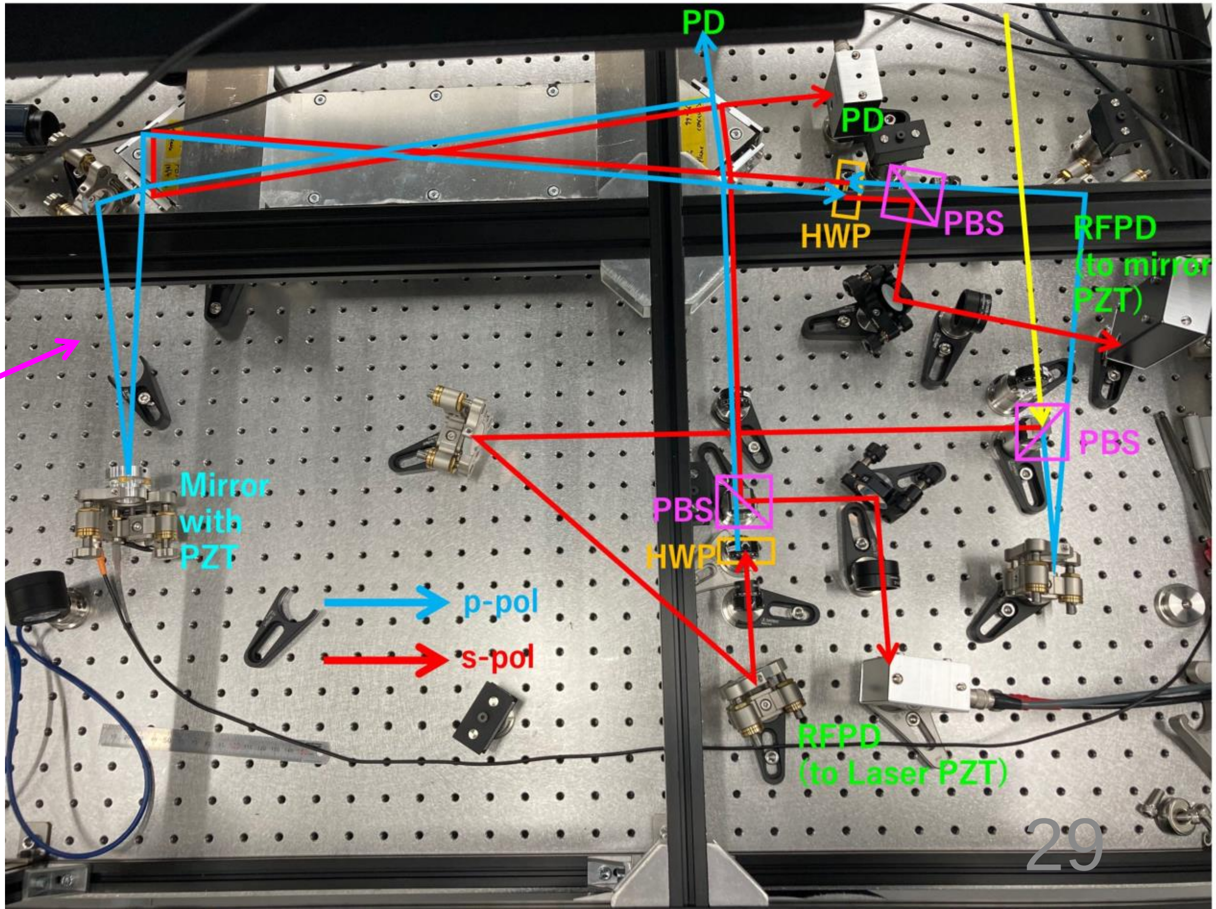
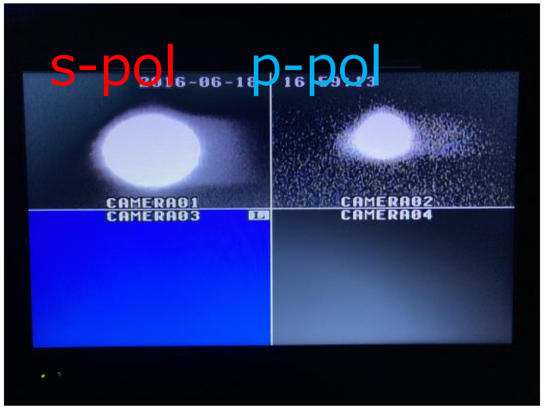


Photo by H. Fujimoto

Simultaneous Resonance Achieved

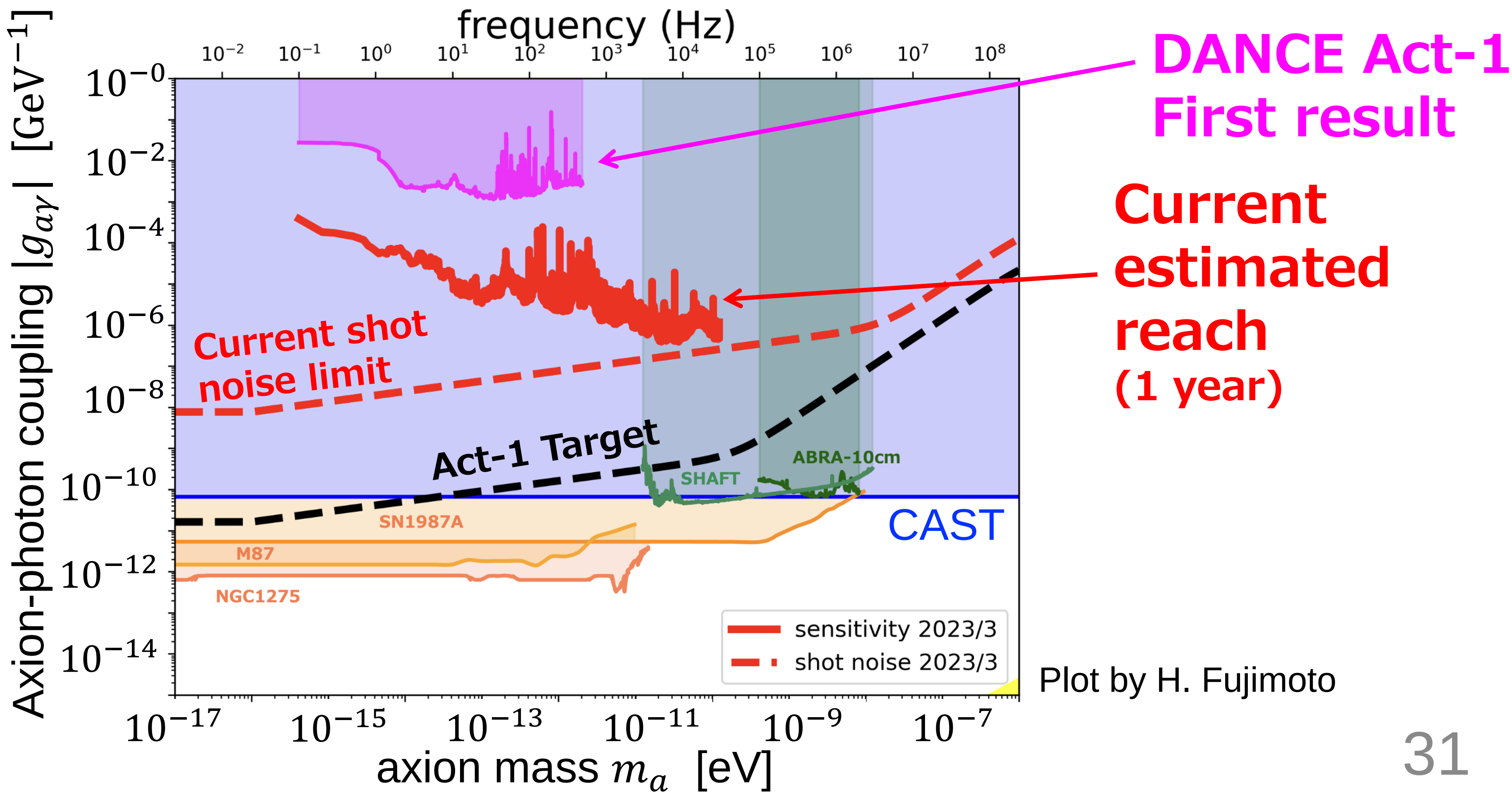
- First demonstration in November 2021
- Finesse reduced due to optical losses in auxiliary cavity



	May 2021	Now (Nov 2022)	Act-1 Target
Round-trip length	1 m	1 m (+0.5 m aux. cavity)	1 m
Input power	242(12) mW (Source: 0.5 W)	21.4(9) mW (Source: 2 W)	1 W
Finesse (for carrier)	$2.85(5)\times10^3$ s-pol	549(3) s-pol, with cavity lock	3×10^3
Finesse (for sidebands)	195(3) p-pol	26.8(2) p-pol, with cavity lock	3×10^3
s/p-pol resonant freq. difference	2.52(2) MHz	~0 Hz with lock (Originally ~92 MHz)	0 Hz

Current Estimated Sensitivity

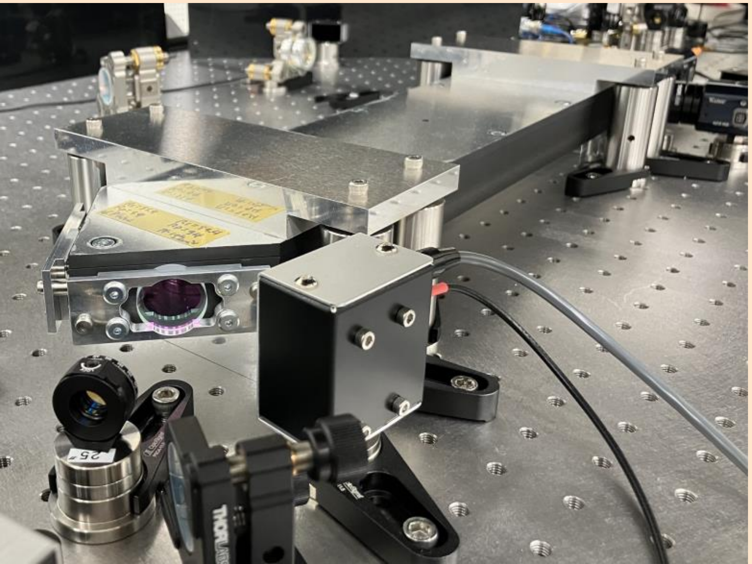
- Improved by more than two orders of magnitude
- Next: Better quality mirrors for improving finesse



Plot by H. Fujimoto

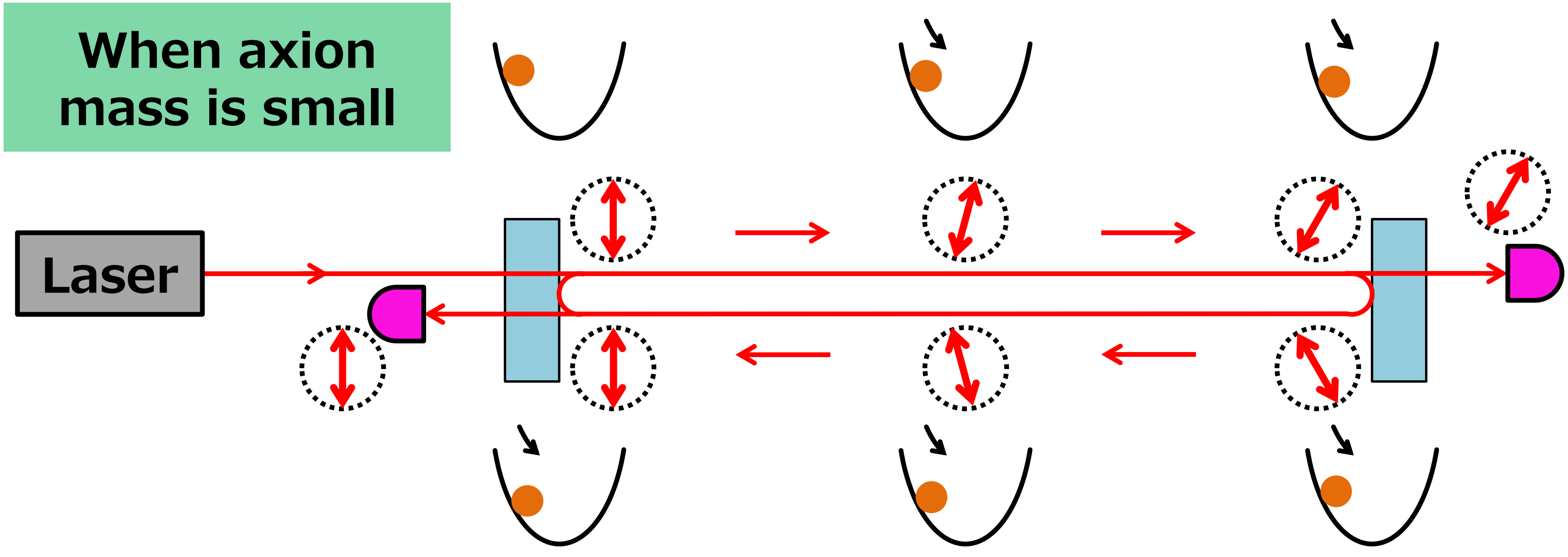
Contents

- Axion dark matter search
  with table-top optical ring cavity
- **Axion dark matter search**
  with gravitational wave detectors
- Vector dark matter search
  with gravitational wave detectors



Linear Cavities for Axion Search

- Polarization flip at mirror reflection can be used to enhance the signal when the round-trip time equals odd-multiples of axion oscillation period
- Long baseline linear cavities in gravitational wave detectors are suitable



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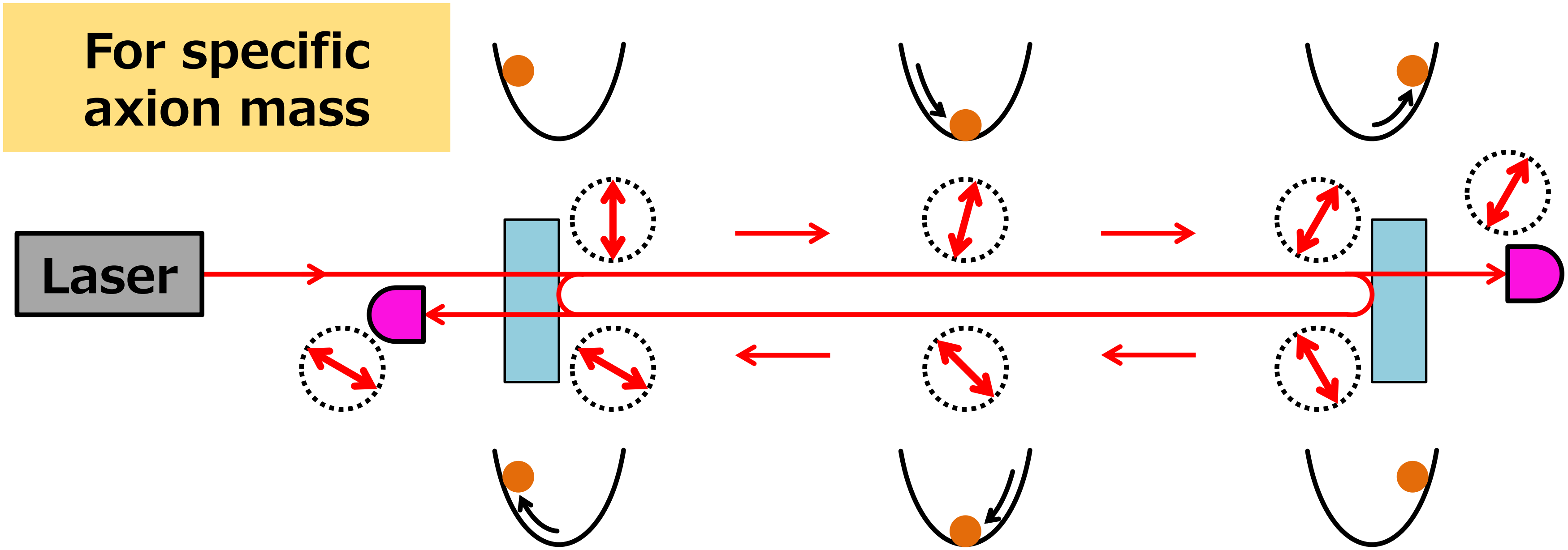
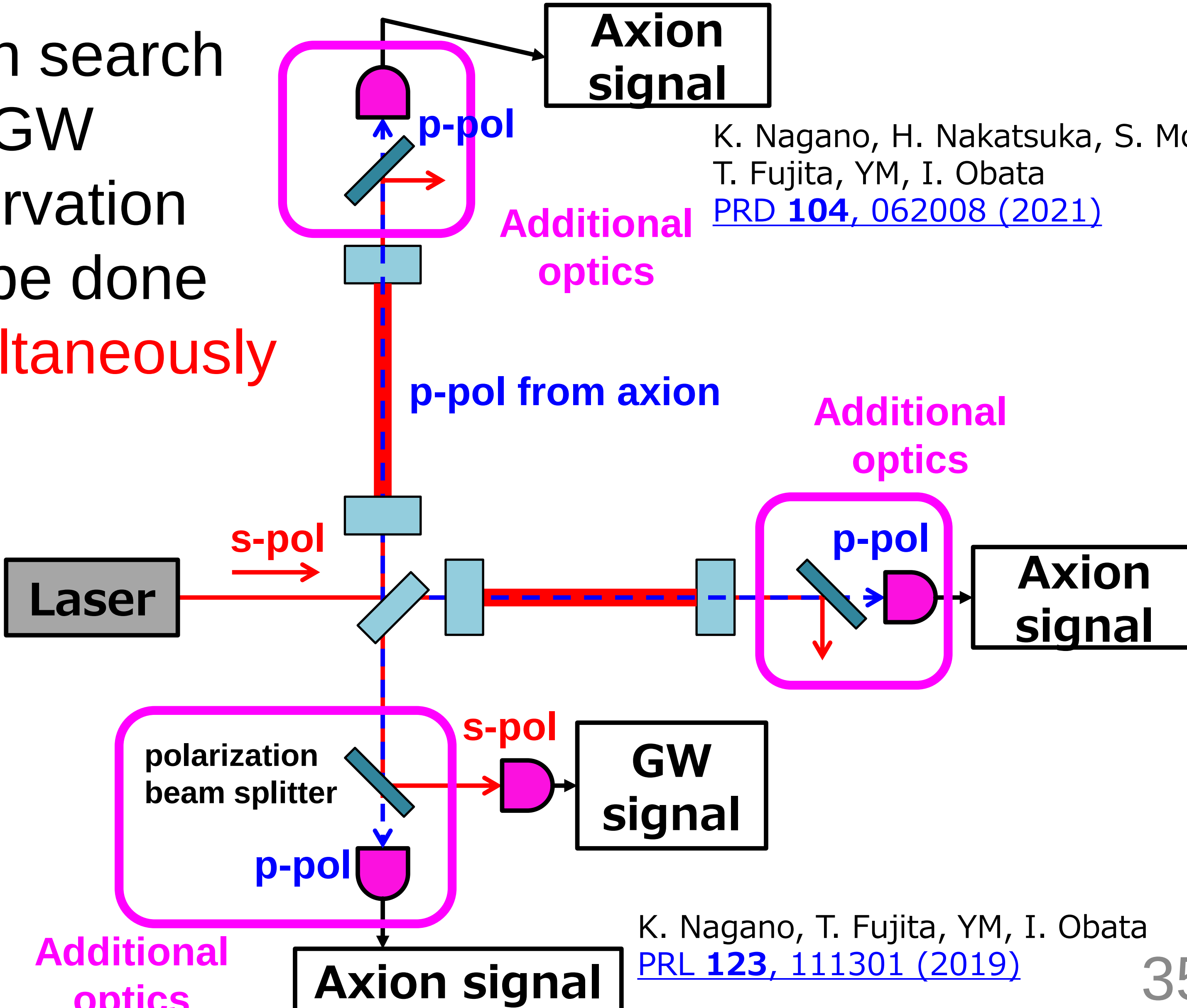
- Polarization flip at mirror reflection can be used to enhance the signal when the round-trip time equals odd-multiples of axion oscillation period
- Long baseline linear cavities in gravitational wave detectors are suitable



Axion Search with GW Detectors

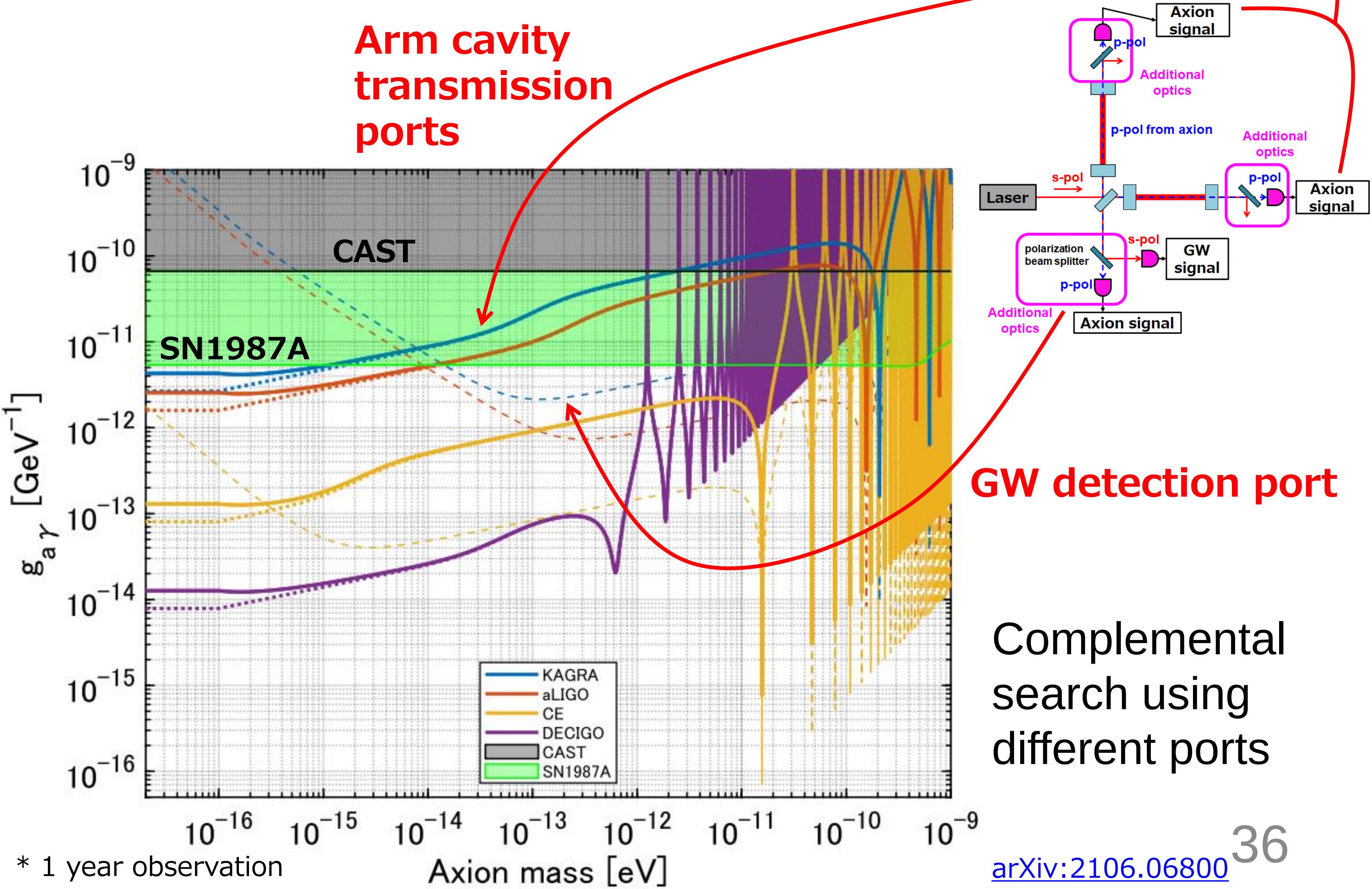
- Axion search and GW observation can be done simultaneously



K. Nagano, H. Nakatsuka, S. Morisaki, T. Fujita, YM, I. Obata
PRD **104**, 062008 (2021)

K. Nagano, T. Fujita, YM, I. Obata
PRL **123**, 111301 (2019)

Axion Sensitivity



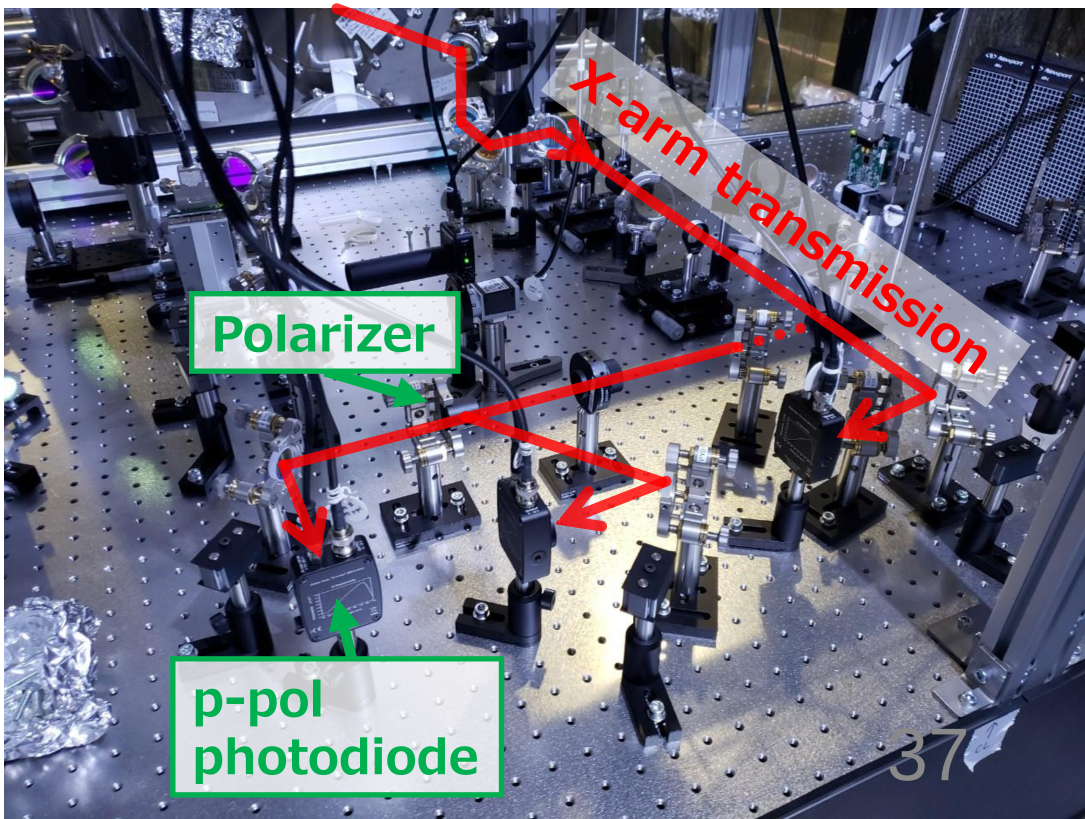
Complemental search using different ports

* 1 year observation

arXiv:2106.06800

Optics for Axion Search Installed

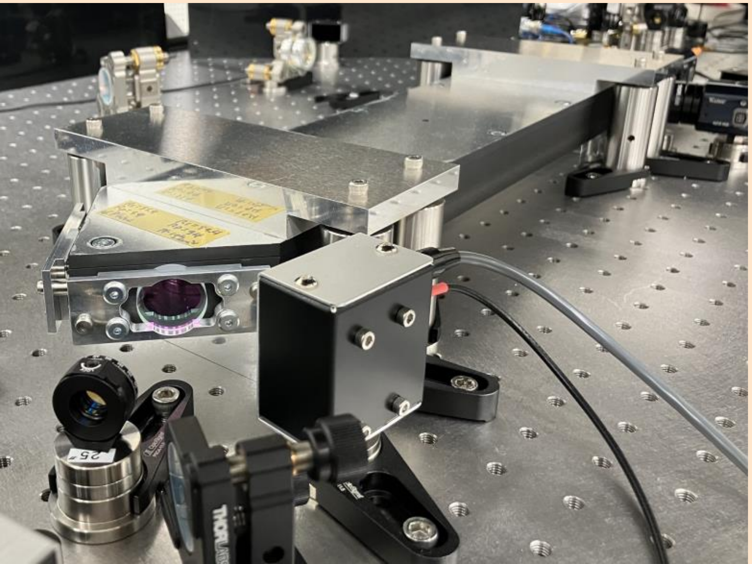
- For KAGRA, polarization optics were installed for X-arm transmission in July 2021 and Y-arm transmission in December 2021
  - Ready to take data in O4 (starting May 2023!)
- For LIGO, auxiliary port of output Faraday isolator can be used (calibration method needs to be developed)



klog #17692

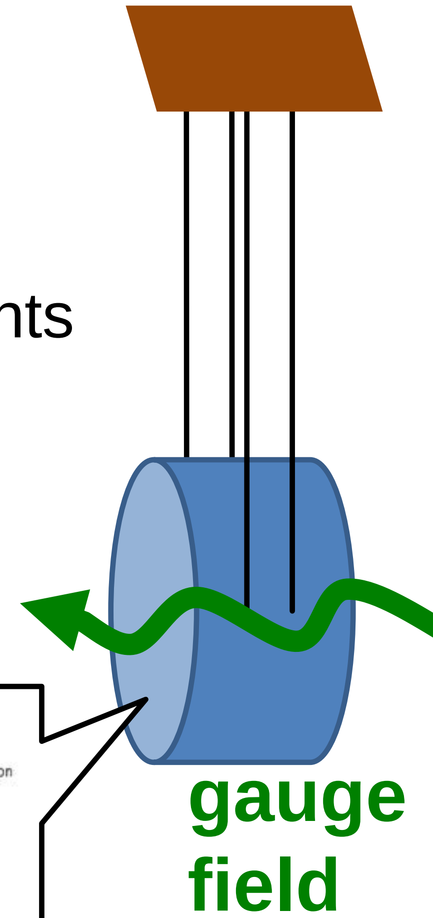
Contents

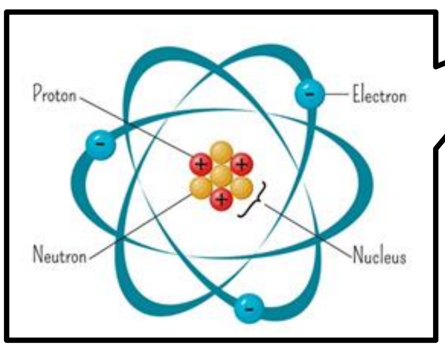
- Axion dark matter search
  with table-top optical ring cavity
- Axion dark matter search
  with gravitational wave detectors
- **Vector dark matter search
  with gravitational wave detectors**



Gauge Boson

- Possible new physics beyond the standard model:
  New gauge symmetry and gauge boson
- New gauge boson can be dark matter
- B-L (baryon minus lepton number)
  - Conserved in the standard model
  - Can be gauged without additional ingredients
  - Equals to the number of neutrons
  - Roughly 0.5 per neutron mass,
    but slightly different between materials
    Fused silica: 0.501
    Sapphire: 0.510
- Gauge boson DM gives oscillating force





Oscillating Force from Gauge Field

- Acceleration of mirrors

charge　　　　　gauge boson mass

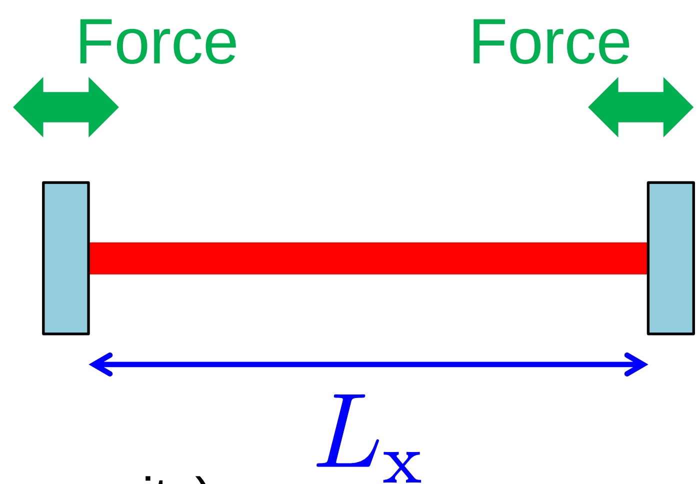
$$\vec{a}(t,\vec{x}) = \epsilon_D e \frac{q_D}{M} \sqrt{2\rho_{DM}}\, \vec{e}_A \sin\left(m_A t - \vec{k}\cdot\vec{x}\right)$$

polarization

coupling　　mirror mass　　DM density　　different phase at different position

(normalized by e)

- Gauge boson mass and coupling can be measured by measuring the oscillating mirror displacement



- Almost no signal for symmetric cavity if cavity length is short
(phase difference is 10^{-5} rad @ 100 Hz for km cavity)
- How about using interferometric GW detectors?

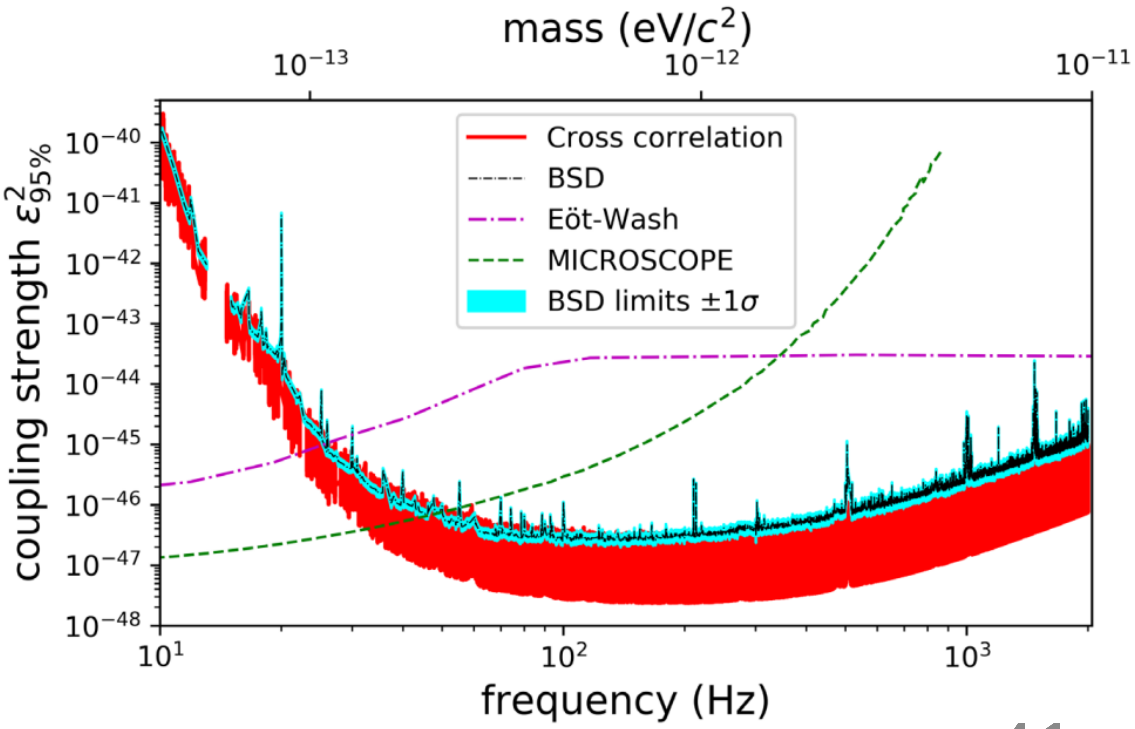
A. Pierce+, PRL **121**, 061102 (2018)

Previous Searches with LIGO/Virgo

- Gauge boson dark matter search with LIGO O1 data and LIGO/Virgo O3 data have been done

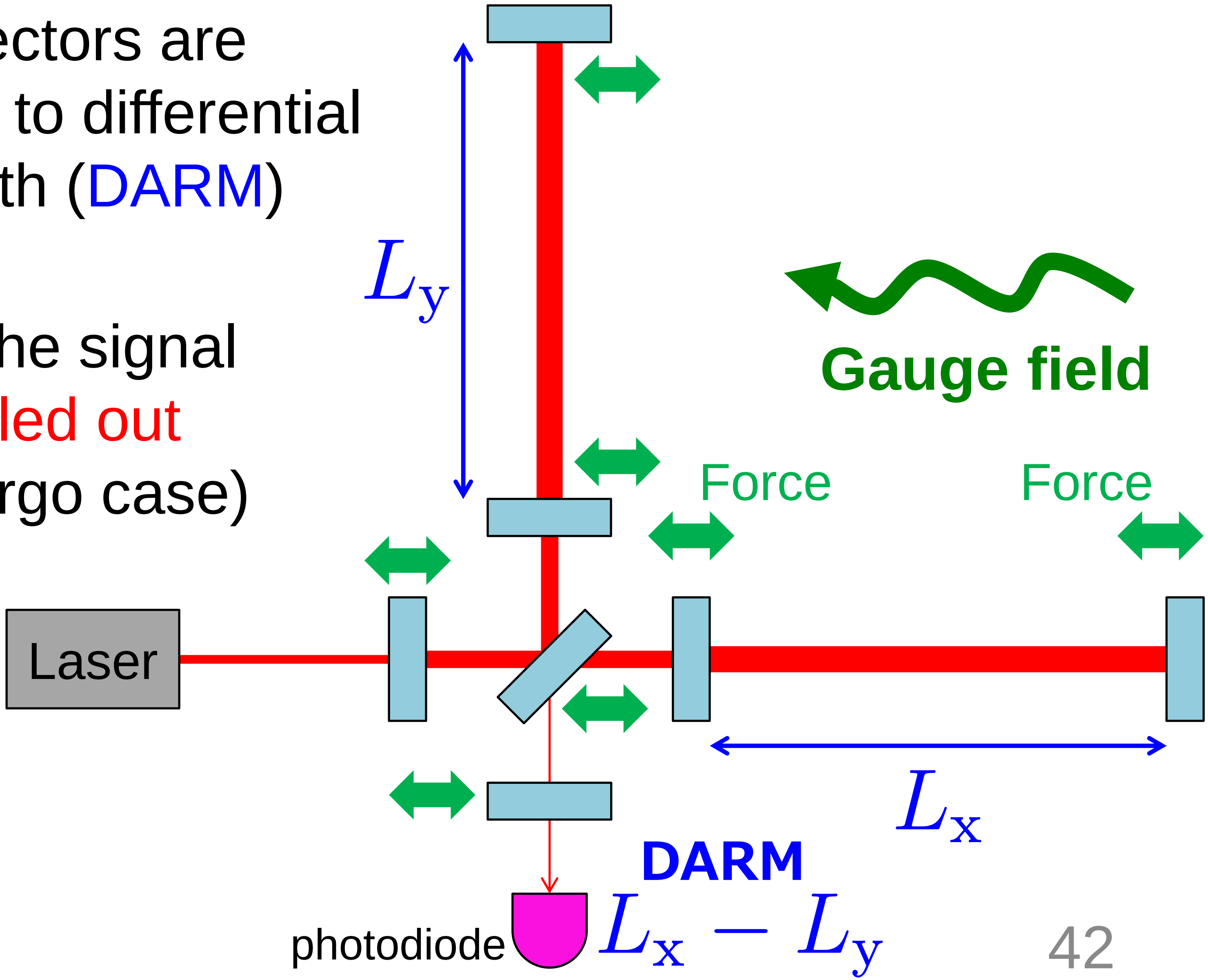
  H-K Guo+, Communications Physics **2**, 155 (2019)
  LIGO, Virgo, KAGRA Collaboration, PRD **105**, 063030 (2022)

- Better constraint than equivalence principle tests
- Even better constraint could be obtained from KAGRA



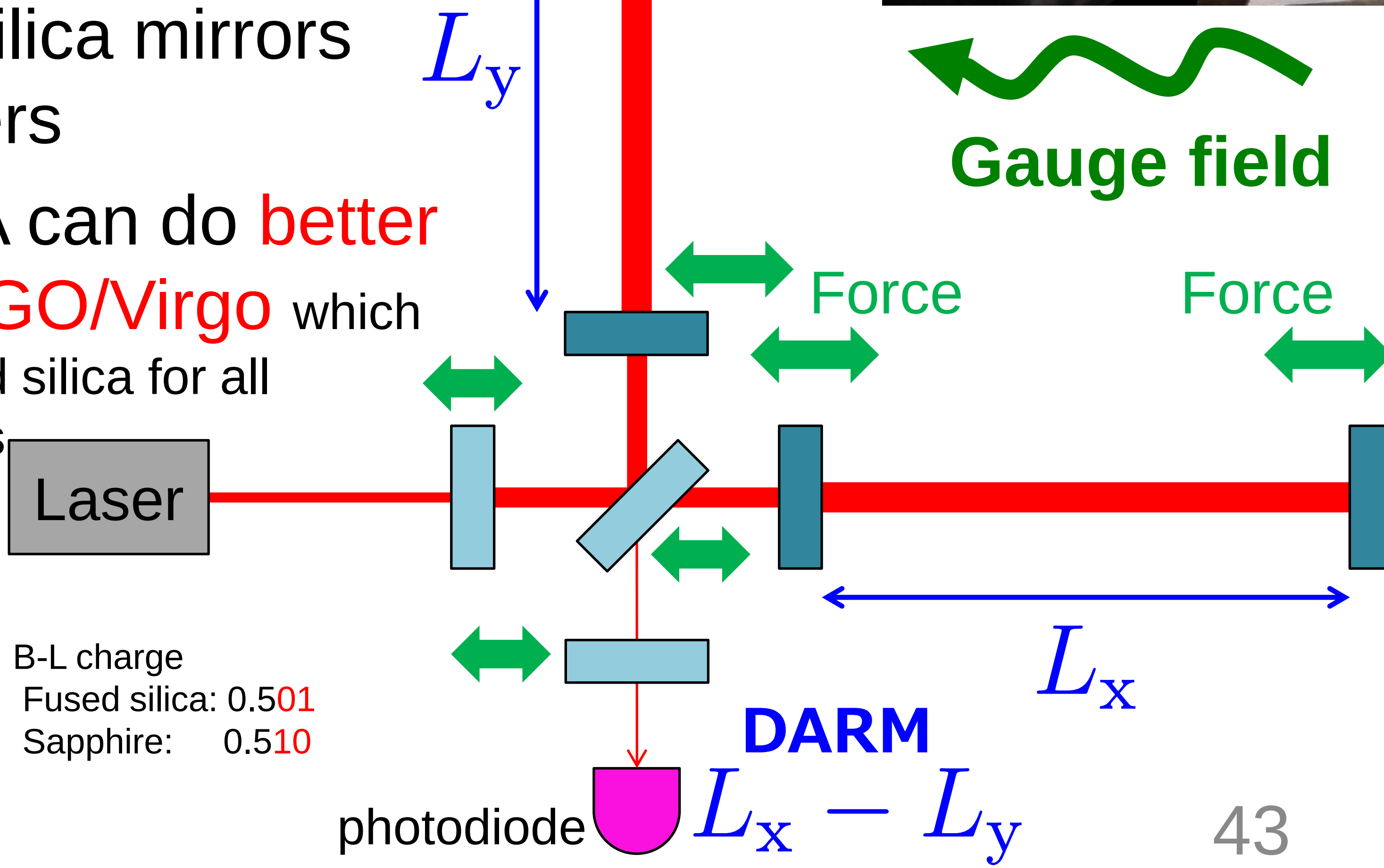
Search with GW Detectors

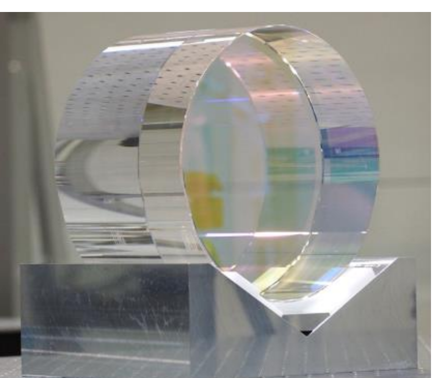
- GW Detectors are sensitive to differential arm length (DARM) change
- Most of the signal is cancelled out (LIGO/Virgo case)



Search with KAGRA

- KAGRA uses cryogenic sapphire mirrors for arm cavities, and fused silica mirrors for others
- KAGRA can do better than LIGO/Virgo which uses fused silica for all the mirrors

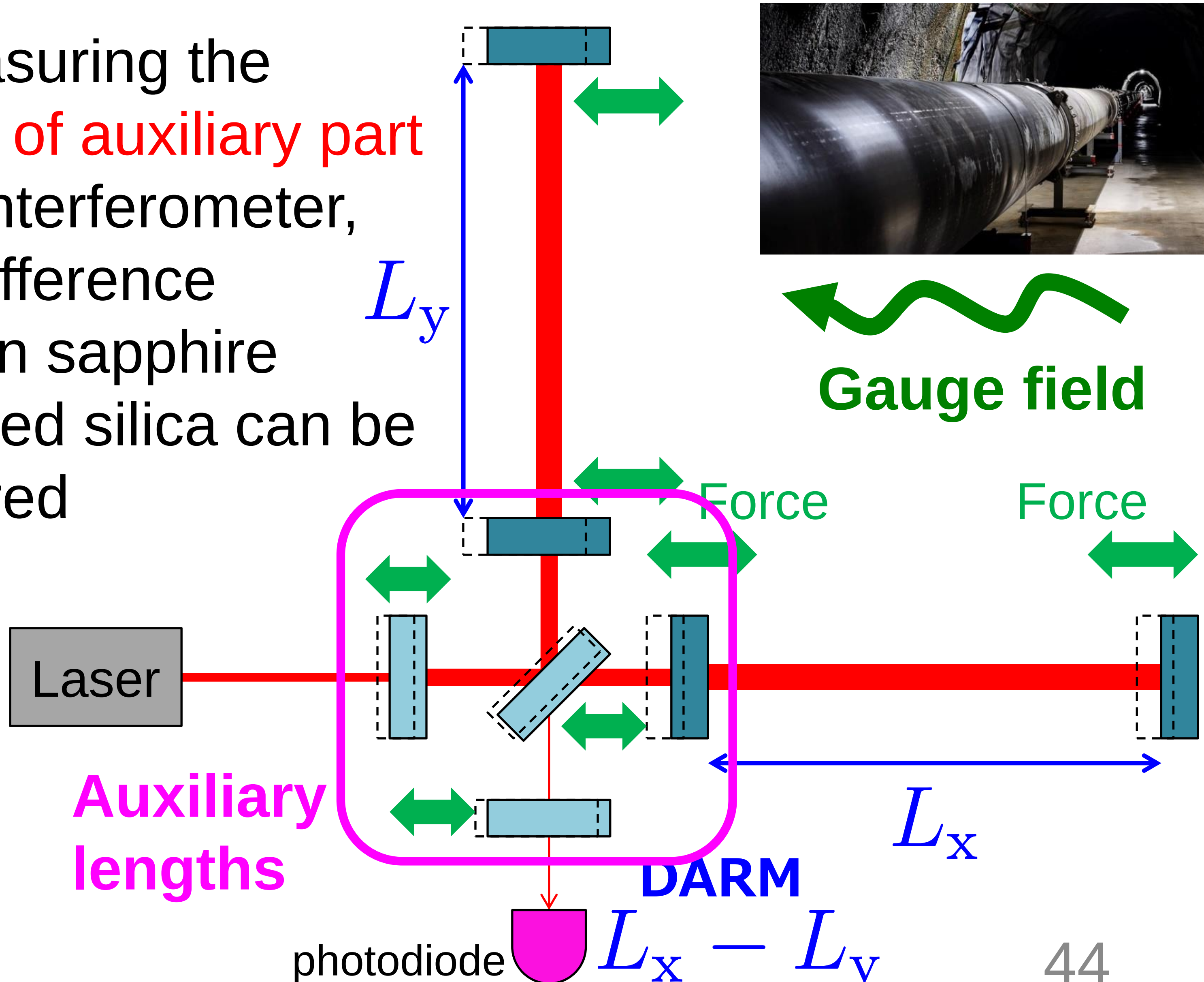




B-L charge
Fused silica: 0.501
Sapphire: 0.510

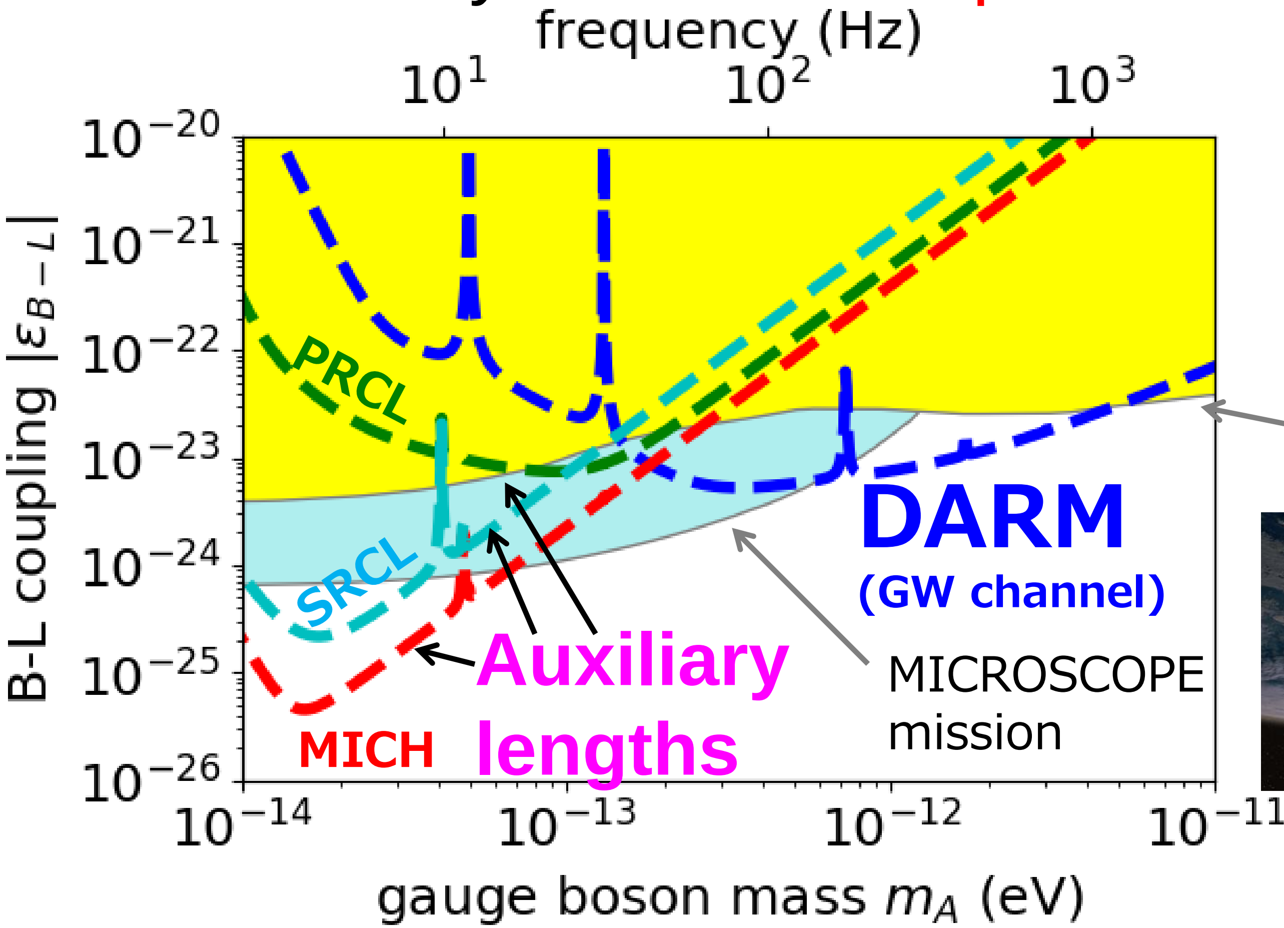
Search with KAGRA

- By measuring the lengths of auxiliary part of the interferometer, force difference between sapphire and fused silica can be measured



KAGRA Gauge Boson Sensitivity

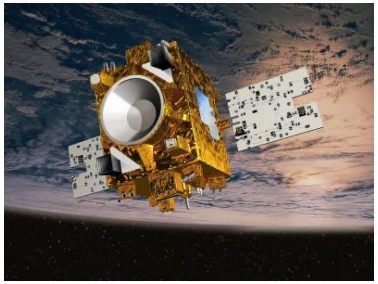
- Auxiliary length channels have better design sensitivity than DARM (GW channel) at low mass range
- Sensitivity better than equivalence principle tests



YM, T. Fujita, S. Morisaki, H. Nakatsuka, I. Obata, PRD 102, 102001 (2020)

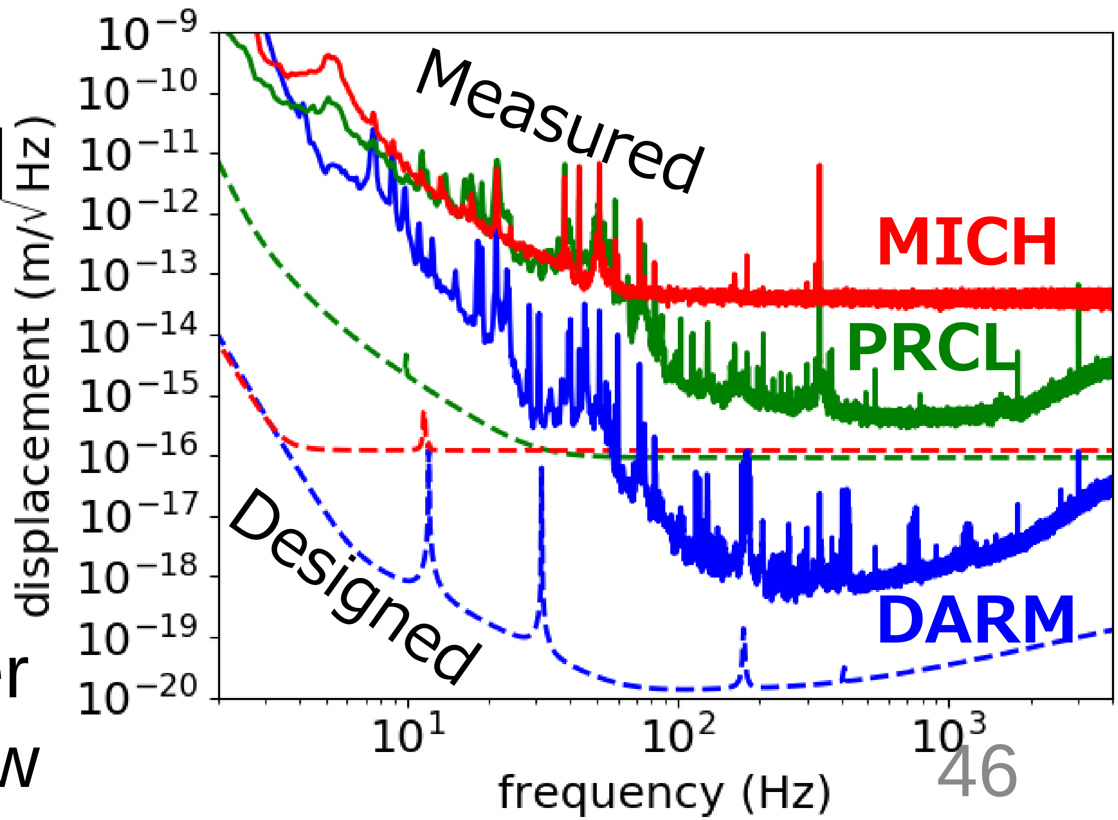
S. Morisaki, T. Fujita, YM, H. Nakatsuka, I. Obata, PRD 103, L051702 (2021)

Eöt-Wash torsion pendulum



KAGRA 2020 Data Analysis

- KAGRA performed joint observing run in April 2020 with GEO600 (O3GK)
- Displacement sensitivity still not good
  ~ 6 orders of magnitude to go at 10 Hz
- Data analysis underway using the same pipeline used for DANCE
  H. Nakatsuka+, [arXiv:2205.02960]
- Results will be available summer 2023 after LVK internal review



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Atsushi Nishizawa
Ippei Obata
Yuka Oshima
Hinata Takidera
Haoyu Wang
… and more to come!

WIAS
KAVLI IPMU
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ICRR Institute for Cosmic Ray Research University of Tokyo
JAXA
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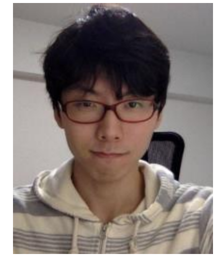
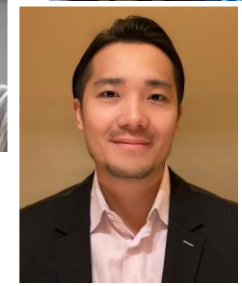
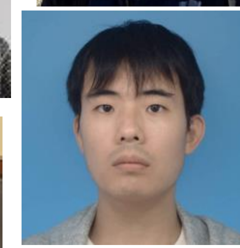
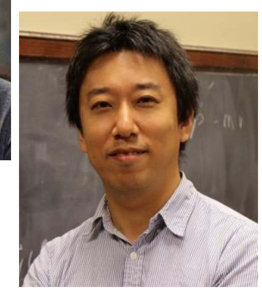
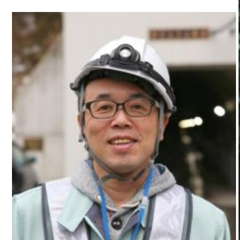
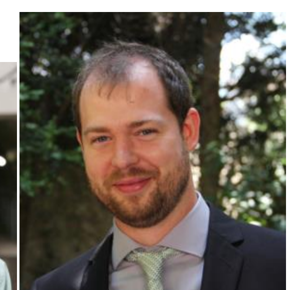
国立天文台 NAOJ National Astronomical Observatory of Japan
UNIVERSITAS STUDIORUM ATHESINA UNITRENTO





Summary

- Laser interferometers open up new possibilities for dark matter search
- **Axion DM search with DANCE**
  - First result from 24-hour data reported
  - Upgrade underway

  Y. Oshima+, arXiv:2303.03594
- **Axion DM search with LIGO-Virgo-KAGRA**
  - Polarization optics installed in KAGRA and LIGO
  - First search to be done in O4 (starting May 2023!)
- **Vector DM search with LIGO-Virgo-KAGRA**
  - Most stringent bound obtained from LIGO-Virgo
  - New search using sapphire mirrors of KAGRA underway

科研費 KAKENHI
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ダークマターの正体は何か？
広大なディスカバリースペースの網羅的研究
What is dark matter? - Comprehensive study of the huge discovery space in dark matter
文部科学省
科学研究費助成事業
学術変革領域研究
（2020–2024）
PRESTO SAKIGAKE
公益財団法人 住友財団
The Sumitomo Foundation

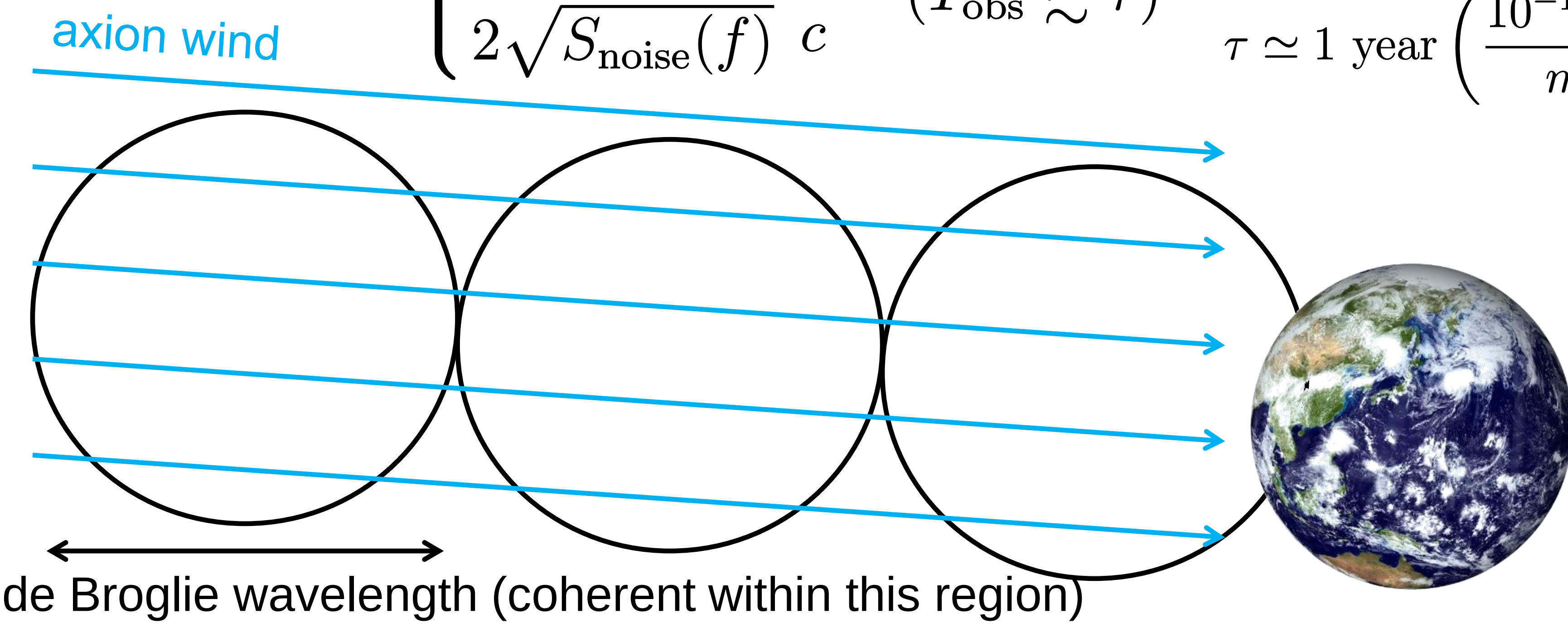
Additional Slides

Coherence Time

- SNR grows with $\sqrt{T_{obs}}$ if integration time is shorter than coherence time
- SNR grows with $(T_{obs})^{1/4}$ if integration time is longer

$$\mathrm{SNR} = \begin{cases} \dfrac{\sqrt{T_{\mathrm{obs}}}}{2\sqrt{S_{\mathrm{noise}}(f)}} \dfrac{\delta c}{c} & (T_{\mathrm{obs}} \lesssim \tau) \\ \dfrac{(T_{\mathrm{obs}}\tau)^{1/4}}{2\sqrt{S_{\mathrm{noise}}(f)}} \dfrac{\delta c}{c} & (T_{\mathrm{obs}} \gtrsim \tau) \end{cases}$$

$$\tau \simeq 1\ \mathrm{year} \left(\frac{10^{-16}\ \mathrm{eV}}{m_a} \right)$$



Freq-Mass-Coherence Time

Frequency	Mass	Coherent Time	Coherent Length
0.1 Hz	4.1e-16 eV	0.32 year	3e12 m
1 Hz	4.1e-15 eV	1e6 sec 12 days	3e11 m
10 Hz	4.1e-14 eV	1.2 days	3e10 m
100 Hz	4.1e-13 eV	2.8 hours	3e9 m
1000 Hz	4.1e-12 eV	17 minutes	3e8 m
10000 Hz	4.1e-11 eV	1.7 minutes	3e7 m